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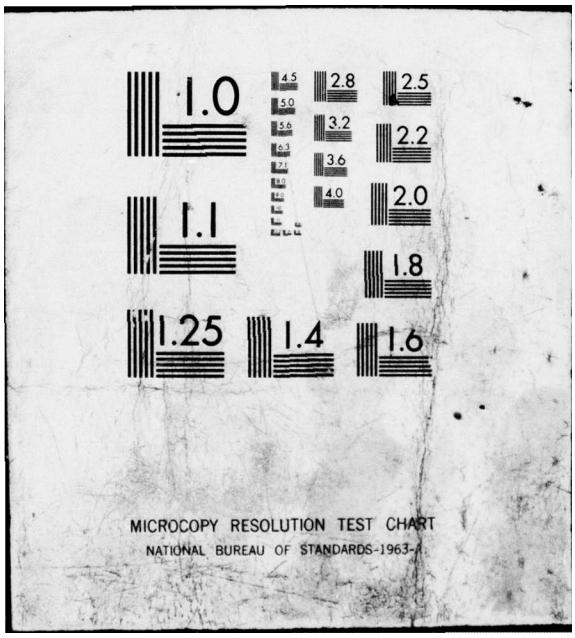
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Final Technical Report
August 1978

SOFT X-RAY PHOTOEMISSION AND CHARGE DEPOSITION NEAR MATE INTERFACES

D. J. Strickland
D. L. Lin
V. W. Pine
W. L. Chadsey

Science Applications Inc.

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This research was sponsored by the Defense Nuclear A under Subtask Z99QAXTA040, Work Unit Code 71 entitl Electron Transport Phenomenology."

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The work reported here comes under two distinct headings: (1) soft x-ray photoemission and (2) charge deposition profiles near irradiated material interfaces. The first electron backscatter results using the SXRP code are reported. Good agreement is obtained with measurements for 5 keV electrons incident on Al. The effect of elastic scattering on the photoemission spectrum is analyzed in further detail following the initial observations made in a previous DNA contract. Additional code validation tests plus the good agreement with measurements for backscatter confirm the —(CONT'D)

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previous result that scattering significantly effects the photoemission spectrum. A simple technique is presented which relates the inner shell ionization IMFP to the corresponding photoionization cross section. The technique is tested on the K and L shells of Aℓ and gives IMFP's which agree well with more rigorous results. The development of an empirical photo-emission code is reported along with spectra and yields for exploding wire sources incident on Aℓ, Au, Ag, and C. Finally charge deposition profiles for several material configurations containing thin foils are obtained with the POEM code and compared with data by Frederickson. In general, good agreement is achieved.

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Section 1

INTRODUCTION AND SUMMARY

The work performed for this contract comes under three distinct headings. They are

(1) User's Guide to X-Ray Dose Enhancement,

(2) Soft X-Ray Photoemission, and

(3) Charge Deposition Profiles Near Irradiated Material Interfaces.

W. L. Chadsey and V. W. Pine participated in the work under headings (1) and (3) while D. J. Strickland and D. L. Lin carried out the work on soft x-ray photoemission. The User's Guide has been separately prepared by W. L. Chadsey. Thus, topics (2) and (3) will be discussed in this report.

The work carried out on soft x-ray photoemission is a continuation of work supported over the past two years by DNA and AFSC.[1-3] During that time,

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[1] D. J. Strickland, "Soft X-Ray Photoemission," RADC-TR-77-252, (July 1977).

[2] D. J. Strickland, D. L. Lin, T. M. Delmer, S. Rodgers, B. Goplen, and W. L. Chadsey, "Soft X-Ray Photoemission, II," DNA Final Report, Submitted October 1977.

[3] W. L. Chadsey, B. L. Beers, V. W. Pine, D. J. Strickland, and C. W. Wilson, "X-Ray Photoemission; X-Ray Dose Enhancement," RADC-TR-77-253, (July, 1977).

the soft x-ray photoemission code SXRP was developed and applied to aluminum (Aℓ), aluminum oxide ($Aℓ_2O_3$), and silicon dioxide (SiO_2) for narrow Gaussian photon sources used to simulate line sources, for blackbody spectra, and for an exploding wire radiator (EWR) spectrum. For line sources, good agreement in total yield was obtained between SXRP calculations and measurements for photoemission from Aℓ over the photon energy range from ∿ 1 to 10 keV. Furthermore, excellent agreement was achieved with Bradford's total yield for a 50 kVp bremsstrahlung source.

A detailed sensitivity analysis was then carried out for Aℓ to determine the sensitivity of both the total and differential photoemission yields to inverse mean free paths (IMFP's) for elastic and inelastic electron interactions, to the initial photoelectron angular distribution, and to variations in photon spectra. An important result of this analysis was that scattering produces a significant effect on the shape of the photoemission spectrum as well as on the total yield. The work discussed in Sections 2 and 3 of this report was prompted by this observation.

Other noteworthy work pertained to a comparison between SXRP calculations and measurements of the photoemission spectrum and total yield for an EWR source incident on Aℓ. The comparison led to a significant readjustment of one set of data obtained with a magnetic spectrometer and to the conclusion that the data from a second experiment involving biased diodes could not be used to obtain a satisfactory emission spectrum. Details appear in the final report

just completed for DNA.[2] This study has ultimately led to a better basic understanding of photoemission for an EWR source and defined the limits of the experiments.

Under this contract, there were several objectives in the continued study of soft x-ray photoemission. They may be generally classified as follows:

(1) to perform backscatter calculations using the SXRP code,

(2) to further analyze the effect of scattering on the photoemission spectrum,

(3) to begin specifying the needed atomic information for new materials [specifically for gold (Au), silver (Ag), and carbon (C)], and

(4) to provide, in a brief period of time an approximate representation of photoemission from Au, Ag, and C for an EWR source.

The backscatter calculations were performed both to increase the applicability of the SXRP code and to

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[2] D. J. Strickland, D. L. Lin, T. M. Delmer, S. Rodgers, B. Goplen, and W. L. Chadsey, "Soft X-Ray Photoemission, II," DNA Final Report, submitted October, 1977.

critically test the elastic scattering cross section representation in the code. The immediate need for this test was to determine if the strength of the scattering effect observed on the photoemission spectrum was real or due in part to an inaccurate representation of the elastic scattering cross section.

Objectives (3) and (4) relate to DNA's SKYNET program, part of which involves the specification of photoemission from various materials irradiated by an EWR source. As a start in examining these materials, the needed atomic information was assembled for performing empirical photoemission calculations based on the Burke model.[4] The empirical approach was initiated to quickly obtain spectral and total yield information for the immediate purpose of supporting SGEMP code validation experiments. Work is now underway which will extend this effort by providing SXRP results for the additional materials.

The results of the work of soft x-ray photoemission are as follows:

(1) Good agreement in backscatter between SXRP results and measurements was obtained for 5 keV electrons incident on Aℓ.

(2) The SXRP code successfully meets the following conditions as applied to photoelectron transport.

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[4] E. A. Burke, "Soft X-Ray Induced Electron Emission," IEEE Trans. Nuc. Sci., NS-24, No. 6, 2505, (1977).

- the bulk solution is independent of scattering,

- the photoelectron spectrum does not change with depth in the absence of scattering.

(3) The above results plus those from previous tests provide convincing evidence that the SXRP code is accurately predicting the effect of scattering on the photoemission spectrum. The effect, which is to reduce the surface spectrum compared to the bulk spectrum in those spectral regions dominated by electrons having lost appreciable energy, is strong even under the simplifying conditions of a uniform isotropic photoelectron source.

(4) The photoemission spectrum measured by Denisov et. al.[5] for 8 keV photons incident on Aℓ closely resembles the bulk spectrum and is thus not a satisfactory representation of the photoemission spectrum.

(5) A technique has been developed which relates the IMFP's for inner shell

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[5] E. P. Denisov, V. N. Shchmelev, A. N. Mezhevich, and M. A. Rumsh, "Analysis of the Energy Structure of X-Ray Photoemission from a Massive Cathode," Fix. Tver. Tela, 6, 2569 (1964) [Sov. Phys.-Solid State 6, 2047 (1965)].

ionization to the corresponding photoionization cross sections. The method has been applied to Aℓ and gives K- and L- shell IMFP's which agree well with independent results. The technique provides an important new capability which will be used to specify currently unavailable inner shell IMFP's for new materials.

(6) Empirically derived photoemission spectra and total yields have been obtained for Aℓ, Au, Ag, and C for four representations of an EWR source. In general, the yield is not particularly sensitive to the choice of photon spectra considered but is to the material. Gold, e.g., photoemits about 20 times stronger than carbon for an EWR source.

For the third task in this contract, the POEM code was used to calculate charge deposition profiles for several material configurations irradiated by a ^{60}Co x-ray source. The calculated profiles were then compared to the measurements by Frederickson.[6] Each configuration consisted of a series of thin foils sandwiched between two thicker slabs of material, each

[6] A. R. Frederickson, "Charge Deposition Profiles Near Irradiated Material Interfaces," IEEE Trans. Nuc. Sci., NS-23, No. 6, 1867, (1976).

greater than an electron range thick for the most energetic photoelectrons. The materials considered were Aℓ, Sn, Ta, and Pb. Results will be presented in Section 6 for six different configurations. To summarize, good agreement was obtained between the calculations and measurements near material interfaces having large atomic number differences. Best agreement was achieved for an Aℓ-Aℓ-Pb configuration with the least satisfactory agreement for an Sn-Sn-Aℓ configuration.

Section 2

FURTHER INVESTIGATION INTO THE EFFECT OF SCATTERING ON THE PHOTOEMISSION SPECTRUM

One of the most important results obtained to date in our soft x-ray photoemission program is that scattering has a pronounced effect on the shape of the photoemission spectrum and its total yield even under the simplifying assumptions that the photoelectron source is uniform and isotropic. Our initial observations were made using an 8 keV photon source incident on Aℓ. This source was chosen since, to our knowledge, it is the only one for which spectral data on photoemission from Aℓ exists over an extended energy range.[5] We found that a serious discrepancy existed between our results and the data. A further problem was the excellent agreement Burke achieved with the data using his empirical model.[4] This model is based on the assumptions that the photoelectron source is uniform and isotropic, and that the electrons lose energy continuously. To account for scattering, Burke introduced a material dependent reflection

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[4] E. A. Burke, "Soft X-Ray Induced Electron Emission," IEEE Trans. Nuc. Sci., NS-24, No. 6, 2505 (1977).

[5] E. P. Denisov, V. N. Shchemelev, A. N. Mezhevich, and M. A. Rumsh, "Analysis of the Energy Structure of X-Ray Photoemission from a Massive Cathode," Fiz. Tver. Tela, 6, 2569 (1964) [Sov. Phys.-Solid State 6, 2047 (1965)].

parameter which typically reduces his total yield by ~50%. The effect of scattering was not, however, taken into account in obtaining the shape of the photoemission spectrum.

Based upon these considerations, a series of SXRP runs was made with a uniform, isotropic photoelectron source and with varying degrees of scattering to see if Burke's spectral shape could be achieved in the limit of zero scattering. The code was, in fact, able to closely reproduce Burke's result which provided an important test on the treatment given to scattering.

Figure 1 summarizes the state of affairs prior to the work performed under this contract. Most of these results have already appeared (in log-log form) in a recent DNA final report.[2] The features peaking at 7.5, 6.2, and 1.4 keV are due to L-photo, K-photo, and KLL Auger electrons respectively. The angular integrated source spectrum used in the calculations is shown in Figure 2. The K- and L- photoelectron features are Gaussian distributions like the 8 keV photon source while the KLL Auger feature is triangular in shape with a width of only 0.2 keV.

To further investigate the observed effect of scattering on the photoemission spectrum, we have

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[2] D. J. Strickland, D. L. Lin, T. M. Delmer, S. Rodgers, B. Goplan, and W. L. Chadsey, "Soft X-Ray Photoemission, II," DNA Final Report, submitted October, 1977.

examined the behavior of our photoelectron solutions as a function of depth under various scattering conditions and (2) have performed backscatter calculations using the SXRP code. Step (1) provides additional means of validating the code while step (2) provides a sensitive test of the applied elastic scattering IMFP.

We can conclude from the backscatter results plus new photoemission results to be presented in the next figure, that the difference between the SXRP spectrum and measurement shown in Figure 1 is not due to a faulty representation of the elastic scattering IMFP. The calculated spectrum obtained prior to this work, is not, however, based on the best value of the scattering parameter η_c (defined in Section 3). A value $\eta_c = 3.2$ should have been used instead of $\eta_c = 1.0$ based on the work of Nigam, et al.[7] which came to our attention while investigating the elastic scattering IMFP in considerably more detail than in earlier work. The backscatter calculations discussed in Section 3 confirm the observation that the larger η_c is more appropriate. We have rerun the SXRP code for $\eta_c = 3.2$ and find that the photoemission spectrum does not significantly change. Spectra will be shown shortly for $\eta_c = 3.2$, $\eta_c = 1.0$ as well as for Moliere's η_c which will demonstrate this fact. We include results for Moliere's η_c since this representation has commonly been used in the past in multiple scattering transport

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[7] B. P. Nigam, M. K. Sundaresan, and Ta-You Wu, "Theory of Multiple Scattering: Second Borm Approximation and Corrections to Moliere's Work," Phys. Rev. 115, 491 (1959).

calculations. It is not the best representation in our formulation, however, based on considerations of low energy electrons interacting with the screened Coulomb field of an atom in a solid. Further details on this subject will be reported under the current DNA contract on soft x-ray photoemission.

The purpose for examining the depth dependence of our photoelectron solutions rests on the following known facts:

(1) In the absence of scattering, the surface solution (photoemission spectrum) must equal the bulk solution and

(2) In the bulk region, the solution cannot be a function of scattering.

We find that the SXRP code satisfies both of these conditions. In the absence of scattering the calculated photoelectron spectrum is constant with depth to within a few percent. With a full scattering treatment, the calculated spectrum comes to within a few percent of the non-scattering result at depths well away from the surface. This is shown in Figure 3 where the photoelectron spectrum is given at depths of 0, 10^{-5}, and 10^{-4} cm for the three choices of η_c discussed above. We see for a given representation of η_c, that the spectrum at 10^{-4} cm is essentially the same as the non-scattering surface spectrum in Figure 1.

We do note some differences in the surface spectrum from one η_c to the next although the scattering effect being addressed in this section is of about the same magnitude for all cases. The best surface spectrum is given for $\eta_c = 3.2$.

To explain the effect shown in Figure 3 (for a given η_c), we simply note that, in the bulk, the spectrum contains more electrons which have lost appreciable energy. For example, the bulk spectrum contains $\sim$4 times more electrons than the surface spectrum at 2 keV, an energy at which the electrons have lost a minimum of 4 keV. In the bulk, many of these electrons originate in the fore hemisphere somewhat closer to the surface and then scatter into the back hemisphere. To make the transition requires a large number of scatterings and is accompanied by significant energy loss. As one approaches the surface, there are fewer electrons available in the fore hemisphere for contributing to the flux in the back hemisphere due to the decrease in available volume. This results in a decrease in the spectrum in its large energy loss regions.

In conclusion, we have offered additional evidence that the SXRP code is correctly predicting the extent to which scattering affects the photoemission spectrum. It has been shown that the code

(1) obtains photoelectron spectra independent of depth in the absence of scattering,

(2) obtains the same bulk solution with or without scattering, and

(3) gives good backscatter results for 5 keV electrons incident on Aℓ (next section).

In addition, it has been previously shown that the code

(4) gives good agreement with photoemission yield data for line sources incident on Aℓ over the photon energy range from ∿ 1 to 10 keV, and

(5) reproduces the shape and magnitude of the empirical photoemission spectrum in the absence of scattering (the total yield for either approach is about twice too large with no scattering corrections).

Based on the evidence presented, neither the data by Denisov, nor the bulk photoelectron spectrum, nor the empirical photoemission spectrum, all of which are very similar, give an accurate representation to the photoemission spectrum. The true spectrum shows greater variations with energy due to the effect of scattering as represented by the SXRP calculations.

Section 3

ELECTRON BACKSCATTER FROM Aℓ FOR 5 keV INCIDENT ELECTRONS

The purpose of the backscatter calculations reported in this section was to test the accuracy of our representation of the elastic scattering IMFP for Aℓ in the low kilovolt region. Our immediate interest was to know if elastic scattering is being adequately described in SXRP calculations for an 8 keV photon source discussed in Section 2.

We use the screened Rutherford formula to specify the elastic scattering IMFP. It's differential form is given by

$$K_{elas}(E,\Theta) = K_{elas}(E)\ p(E,\Theta)\ \ cm^{-1}sr^{-1} \qquad (1)$$

where

$$p(E,\Theta) = \frac{\eta(\eta+1)}{\pi}\ \frac{1}{(1-\cos+2\eta)^2} \qquad (2)$$

$p(E,\Theta)$ is the normalized Rutherford formula and η is the screening parameter which contains the energy dependence in p. $K_{elas}(E)$ has the form

$$K_{elas}(E) = n\ \frac{Z^2e^4}{v^2p^2}\ \frac{\pi}{\eta(\eta+1)} \qquad (3)$$

where v and p are respectively the electron velocity and momentum corresponding to E, Z is the atomic number, and η is the material number density. The screening parameter η may be represented in the form

$$\eta(E) = 4.3\, Z^{2/3}\, \frac{\eta_c}{E} \tag{4}$$

where E is in eV and η_c is a function of the screening potential. The screening potential in conducting solids can be adequately represented by a single Yukawa potential which leads to constant η_c.[7] We will consider the values $\eta_c = 1$ and 3.2. We have previously used $\eta_c = 1$ since values close to unity have been applied before.[8,9] The value 3.2 appears more appropriate since it is based on considerations of the screening potential itself.[7]

Moliere has presented an energy dependent form for η_c which has had wide application to multiple scattering calculations at high energy. It is given by

$$\eta_c = 1.13 + 3.76 \left(\frac{Z}{137 \frac{v}{c}}\right)^2 \tag{5}$$

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[7] B. P. Nigam, M. K. Sundaresan, and Ta-You Wu, "Theory of Multiple Scattering: Second Born Approximation and Corrections to Moliere's Work," Phys. Rev. 115, 491 (1959).

[8] M. J. Berger, S. M. Seltzer, and K. Maeda, "Energy Deposition by Aurola Electrons in the Atmosphere," J. Atmos. Terr. Phys., 32, 1015 (1970).

[9] H. E. Bishop, "Electron Scattering in Thick Targets," Brit. J. Appl. Phys., 18, 703, (1967).

Results will also be presented for this η_c primarily to show the sensitivity of backscatter to different representations of $K_{elas}(E,\theta)$. We do not suggest the use of Moliere's η_c at low energies in our formulation which is based on a single scattering description, i.e. where each type of electron event is modelled in detail.

For the three η_c's just discussed, Figure 4 shows the resulting screening paramenters and the elastic IMFP's. The energy range extends down to the lower limit treated in the present calculations.

The point of the present backscatter calculations, as stated above, was to test the values of $K_{elas}(E)$ and $p(E,\theta)$ which have been used in photoemission calculations. In particular, we are interested in initial electron energies in the vicinity of 5 keV which dominate the electron source spectrum for the 8 keV photon source we have been using to investigate the effect of scattering on the photoemission spectrum. The incident electron flux we have chosen is Gaussion in both E and μ and peaks at 5 keV. Designating the incident flux by $\phi_o(E,\mu)$ ($\phi_o=\phi(z=0,E,\mu>o)$ it is given by

$$\phi_o(E,\mu) = f(E)g(\mu) \quad \text{el/cm}^2\text{-sec-eV-sr} \qquad (6)$$

with

$$f(E) = \frac{4}{\sqrt{\pi}} \exp\left[-\left(\frac{E-5}{\cdot 25}\right)^2\right] \qquad (7)$$

and

$$g(\mu) = \frac{10}{\sqrt{\pi}} \exp\left[-\left(\frac{1-\mu}{\cdot 2}\right)^2\right] \qquad (8)$$

The functions f and g are both normalized over their respective ranges. The choice of width for f(E) is not critical since the backscatter coefficient or yield for Aℓ varies slowly with energy at 5 keV. The choice of width for g(μ) was dictated by the number of μ points used in the calculations (20). g(μ) does not give an effective representation of normal incidence but this is not necessary for the testing of $K_{elas}(E)$ and $p(E,\Theta)$ since we are able to adjust our backscatter yields for comparison with available data which are typically taken at normal incidence. The adjustment factor is simply

$$r = \frac{Y(\mu=1)}{\int_0^1 g(\mu)Y(\mu)d\mu} \qquad (9)$$

where Y(μ) is the backscatter yield as a function of μ, the cosine of the incidence angle. Y (μ)/Y(1) is shown in Figure 5 along with g(μ). The shape of Y(μ) was taken from Bishop for 10 keV electrons incident on Aℓ.[9] We expect the shape of Y(μ) at 5 keV to be similar based on, e.g., the work of Bishop[9] and

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[9] H. E. Bishop, "Electron Scattering in Thick Targets," Brit. J. Appl. Phys., 18, 703, (1967).

Darlington and Cosslett.[10] The resulting value of r is .75 which will be applied shortly.

For the incident flux given by equations (6)-(8), Figure 6 shows the differential and cumulative backscatter yields for four different representations of $K_{elas}(E,\Theta)$. Table 1 gives the total yields as well as adjusted values for normal incidence based on r = .75. Three of the cases considered refer to equations (1)-(3) with $\eta_c = 1$, $\eta_c = 3.2$, and Moliere's η_c. The fourth is given by Moliere's η_c with $K_{elas}(E)$ scaled down by a factor of 2. This case is included just to show the sensitivity of backscatter to the magnitude of $K_{elas}(E)$ holding the mean angle of scattering fixed or equivalently holding η fixed. For what we believe to be the best representation of K_{elas}, i.e., for $\eta_c = 3.2$, the backscatter yield is 21% which when corrected for non-normal incidence is 16%. This is in excellent agreement with measurements of this quantity.[4] The Moliere representation also gives excellent agreement with the data as well as with the differential yield above 2 keV for $\eta_c = 3.2$. The agreement between the two representations is due to the intersection of their η's near 4 keV which is in the region primarily responsible for the backscatter characteristics. We see that the yield is quite sensitive to the scaling of $K_{elas}(E)$. There is nearly a one-to-one correspondence between the yield and scaling factor for the given material and given energy. There is much less sensitivity, however, if instead of charging the

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[4] E. A. Burke, "Soft X-Ray Induced Electron Emission," IEEE Trans. Nuc. Sci., NS-24, No. 6, 2505, (1977).

[10] E. H. Darlington and V. E. Cosslett, "Backscattering of 0.5 - 10 keV Electrons from Solid Targets," J. Phys. D: Appl. Phys., 5, 156, (1972).

overall magnitude of K_{elas} by scaling, the magnitude is changed using a different dependence in η. For example, $K_{elas}(E)$ for $\eta_c = 1$ is significantly different than for Moliere's η_c (Figure 4) although the corresponding backscatter yields are similar. This results from nearly compensating effects due to the mean angle of scattering changing with η.

TABLE 1. BACKSCATTER YIELDS Y_{BS} FOR THE INCIDENT FLUX GIVEN BY EQUATIONS (6)-(8) AND THEIR CORRESPONDING YIELDS rY_{BS} ADJUSTED TO NORMAL INCIDENCE. THE CONSTANT C IS THE FACTOR SCALING $K_{elas}(E)$. THE MEASURED VALUE FOR COMPARISON WITH rY_{BS} is $\sim$ .17.

η_c	C	Y_{BS} (el/el)	rY_{BS} (el/el)
3.2	1	0.21	0.16
1	1	0.29	0.22
Moliere	1	0.21	0.16
"	0.5	0.12	0.09

Section 4

A SIMPLE TECHNIQUE FOR OBTAINING INNER SHELL IMFP'S FROM PHOTOABSORPTION COEFFICIENTS

The SXRP code requires IMFP's for the various types of electron interactions in solids. An important interaction, especially for high Z materials, is inner shell ionization. There is currently little information available on IMFP's for this process. The choice of materials previously investigated using the SXRP code (Aℓ, $Aℓ_2O_3$, and SiO_2) was dictated by the availability of such information.[11,12] There is now interest in extending the applicability of the SXRP code to new materials, in particular to Au, Ag, and C for which the needed IMFP information is not available. This has led us to develop the technique reported herein. Following a brief discussion of the method, IMFP's will be presented for Aℓ and compared with independent results.

We designate the inner shell DIMFP by $K_i(E',W)(cm^{-1}\text{-}ev^{-1})$ for the i^{th} shell where E′ is the incident electron energy and W is the energy loss.

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[11] J. C. Ashley, C. J. Tung, V. E. Anderson and R. H. Ritchie, "Inverse Mean Free Path, Stopping Power, CSDA Range, and straggling in Aluminum and Aluminum Oxide for Electrons of Energy ≤10 keV," Report No. AFCRL-TR-75-0583 (December 1975).

[12] C. J. Tung, J. C. Ashley, V. E. Anderson, and R. H. Ritchie, "Inverse Mean Free Path, Stopping Power, CSDA Range, and Straggling in Silicon and Silicon Dioxide for Electrons of Energy ≤10 keV," Report No. RADC-TR-76-125 (1976).

W may be given by either $E' - E$ or by $E_s + B_i$ where E is the outgoing energy of the degraded incident electron, E_s is the energy of the secondary electron, and B_i is the binding energy for the i^{th} shell. The relationship we have derived between K_i and $\mu_i(h\nu)$, the photoabsorption coefficient is the following:

$$K_i(E',W) = C \frac{\mu_i(h\nu)}{E'h\nu} \ln \frac{(h\nu)^2+4E'(h\nu-B_i)}{(h\nu)^2} \quad (10)$$

where the energy loss W is equated to $h\nu$ and the constant C is

$$C = \frac{R_y}{\pi\alpha} \quad . \quad (11)$$

R_y is the Rydberg energy (12.6 eV) and α is the fine structure constant (1/137). The energies E' and $h\nu$ are in the units of electron volts.

Expression (10) is based upon

(1) the Born approximation,

(2) using the dipole term in the expansion of the generalized oscillator strength, and

(3) allowing the momentum of the incident electron to be shared by the two outgoing electrons.

To obtain the IMFP $K_i(E')$, equation (10) is integrated over $h\nu$ from B_i to E'. Most of the contribution to $K_i(E')$ comes from $h\nu$ values near B_i since $h\nu$ is smallest and μ_i is largest there. Figure 7 shows $K_i(E')$ for the K and L shells of Aℓ as well as more exact calculations by Ashley, et al.[11] The absorption coefficients μ_K and μ_L used in equation (10) are shown in Figure 8. At energies above the peak in either IMFP (K or L), excellent agreement is achieved with the more exact values. The overshoot below the peaks can be attributed to the Born approximation which always leads to an overestimate of K_i near threshold. We are now in the process of incorporating a correction factor to improve the results in this region.

The agreement shown in Figure 7 is significant to our overall program in soft x-ray photoemission since it suggests that equation (10) can provide reasonable values of the inner shell IMFP for new materials. The needed quantities, $\mu_i(h\nu)$ and B_i are readily available for most materials. In actual practice, however, the total absorption coefficient $\mu(h\nu)$ is usually the available experimental quantity which must be decomposed into its separate parts. One can, however, do this with sufficient accuracy to provide a good approximate representation of the needed IMFP's.

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[11] J. C. Ashley, C. J. Tung, V. E. Anderson, and R. H. Ritchie, "Inverse Mean Free Path, Stopping Power, CSDA Range, and Straggling in Aluminum and Aluminum Oxide for Electrons of Energy ≤10 keV," Report No. AFCRL-TR-75-0583 (December 1975).

Section 5

A CODE FOR EMPIRICALLY PREDICTING PHOTOEMISSION AND ITS APPLICATION TO PHOTOEMISSION FROM Aℓ, Au, Ag, AND C FOR AN EXPLODING WIRE RADIATION SOURCE

In this section, we discuss a code developed under this contract which is based on Burke's empirical model for photoemission.[4] The code calculates the total and differential photoemission yields for an arbitrary photon spectrum. This can be a set of lines, a continuum, or a combination of both.

The primary motivation for generating the code came from SAI's involvement in SGEMP code validation experiments sponsored by DNA. We were seeking an approximate representation of the photoemission spectrum for an EWR source incident on Au, Ag and C. An additional reason for generating the code was to provide information against which SXRP results could be compared. This is particularly useful when there are no other data available for comparison.

The needed relationships for line and continuum sources are presented next. This is

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[4] E. A. Burke, "Soft X-Ray Induces Electron Emission," IEEE Trans. Nuc. Sci., NS-24, No. 6, 2505 (1977).

followed by a presentation of the required input data for the materials Aℓ, Au, Ag, and C. Empirical photoemission results are then given for four representations of an EWR source which show both the general features of photoemission for the various materials and the sensitivity of the yield and spectral shape to changes in the photon source spectrum.

The basic empirical expressions for the transport of electrons produced by monoenergetic photon sources were originally presented by Schaefer.[13] They are based on the assumptions of:

- depth independent isotropic electron production,
- continuous energy loss, and
- straight line paths of electrons from creation to escape.

The effect of scattering was accounted for by assuming the penetration range was half the total range. Burke has recently modified Schaefer's expressions by using a reflection coefficient to account for scattering rather than halving the range. The expressions below will contain Burke's coefficient. We begin by giving the yield expressions for line source [Equations (12)-(15)]. The corresponding expressions for a

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[13] R. R. Schaefer, J. Appl. Phys. 44, 152 (1973).

continuum source will follow [Equations (17)-(20)]. The total yield for a given photoelectron component is:

$$Y_i(h\nu) = \frac{(1-\beta)}{4} \mu_i(h\nu) R(h\nu - B_i) \quad \text{electrons/photon} \quad (12)$$

where μ_i is the photoabsorption coefficient (cm^{-1}) for producing the i^{th} type of photoelectron and R is the continuous slowing down approximation (CSDA) range (cm). Its argument contains the starting kinetic energy of the photoelectron where B_i is the i^{th} binding energy. The parameter β is Burke's reflection coefficient. The assumed values of β will be presented later.

The corresponding expression for the j^{th} Auger electron for the i^{th} type of vacancy is:

$$Y^A_{ij}(h\nu) = \frac{(1-\beta)}{4} w_{ij} \mu_i(h\nu) R(E^A_{ij}) \text{ electrons/photon} \quad (13)$$

where w_{ij} is the ij^{th} Auger yield and E^A_{ij} is the starting energy of the ij^{th} Auger electron.

The corresponding differential yields for photo- and Auger electrons are:

$$\frac{dY_i(h\nu,E)}{dE} = \frac{(1-\beta)}{4} \frac{1}{S(E)} \mu_i(h\nu) \quad (14)$$

electrons/photon-keV

$(E \leq h\nu - B_i)$

and

$$\frac{dY_{ij}^{A}(h\nu,E)}{dE} = \frac{(1-\beta)}{4} \frac{w_{ij}}{S(E)} \mu_i(h\nu) \qquad (15)$$

electrons/photon-keV

$(E \leq E_{ij}^{A})$

where S(E) is the stopping power (eV/cm).

Integration of the differential yields from the starting energy to zero gives the total yields since the CSDA range is related to the stopping power by

$$R(E_{max}) = \int_{o}^{E_{max}} \frac{dE}{S(E)} \qquad (16)$$

For a continuum source, the above yields become

$$Y_i = \frac{(1-\beta)}{4} \frac{1}{F} \int_{B_i}^{h\nu_{max}} I(h\nu)\mu_i(h\nu)R(h\nu-B_i)dh\nu \qquad (17)$$

$$Y_{ij}^{A} = \frac{(1-\beta)}{4} \frac{w_{ij}}{F} R(E_{ij}^{A}) \int_{B_i}^{h\nu_{max}} I(h\nu)\mu_i(h\nu)dh\nu \qquad (18)$$

$$\frac{dY_i(E)}{dE} = \frac{(1-\beta)}{4} \frac{1}{FS(E)} \int_{E+B_i}^{h\nu_{max}} I(h\nu)\mu_i(h\nu)dh\nu \quad (19)$$

$$\frac{dY^A_{ij}(E)}{dE} = \frac{(1-\beta)}{4} \frac{w_{ij}}{FS(E)} \int_{E+B_i}^{h\nu_{max}} I(h\nu)\mu_i(h\nu)dh\nu \quad (20)$$

where I is the photon spectrum (photons/cm^2-sec-eV) and F is the total number of photons/cm^2-sec given by

$$F = \int_{h\nu_{min}}^{h\nu_{max}} I(h\nu)dh\nu \quad . \quad (21)$$

We have programmed all of the above yield expressions into the empirical code. As noted earlier, the code can treat either one or more lines, a pure continuum, or a combination of lines plus continuum. Total and differential yields are printed out for each type of photoelectron and Auger electron considered, along with their sums.

The basic data needed for the empirical calculations are:

- photoabsorption coefficients $\mu_i(h\nu)$,
- CSDA range R(E),
- Stopping power S(E),
- Auger energies E^A_{ij} and yields w^A_{ij}, and
- Reflection coefficient β.

We begin by presenting Table 2 which contains the shells and their binding energies, as modeled in this work for Au, Ag, Aℓ, and C. We do not include the K shell of Ag nor the K or L shells of Au since their binding energies lie above the energy range of interest for EWR sources.

Figures 9-12 show the photoabsorption cross sections (cm^2) for the four materials. In each figure, individual shell contributions are shown along with their sum. The total cross section is based on information from Biggs and Lighthill,[14] Storm and Israel,[15] and Hubbell.[16] The extrapolations of individual cross sections were made by the authors.

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[14] F. Biggs and R. Lighthill, Report No. SC-RR-71-0507, Sandia Laboratories, (1971).

[15] E. Storm and H. I. Israel, Report Number LA-3753, Los Alamos Scientific Laboratory, (1967).

[16] J. H. Hubbell, Atomic Data, 3, 241, (1971).

TABLE 2. ATOMIC SHELLS TREATED IN THIS WORK AND THEIR BINDING ENERGIES IN keV

MATERIAL	SHELLS	BINDING ENERGY
Au	M_1	3.425
	M_2	3.150
	M_3	2.743
	M_4	2.291
	M_5	2.206
	N	0.083
	O_{23}	0.052
	O_{45}	0.008
Ag	L	3.351
	M	0.367
	N	0.005
Aℓ	K	1.560
	L	0.073
	M	0.005
C	K	0.284
	L	0.013

Stopping powers and ranges for the four materials are given in Figures 13-16. The information comes from Ashley, et. al.[17] (Au and Ag), Ashley, et. al.[11] (Aℓ), and the Bethe formula [18] (C).

Our model for the relevant Auger spectroscopy is given in Table 3. The model serves its one basic requirement – the proper accounting of the bulk of potential energy of inner shell vacancies. For most transitions shown in the table, several Auger lines exist although one will usually be dominant.[19] The energy shown is the energy of the dominant line.

We have given special attention to gold since it is an important material in satellite systems and since its Auger spectrocopy is complicated. If one refers to, e.g., the handbook of Auger electron spectroscopy,[19] it will be seen that basically three pairs of lines exist between 1.5 and 2.1 keV along with two strong lines at 0.069 and 0.239 keV. There are no other important Auger features in the soft x-ray

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[11] C. J. Tung, J. C. Ashley, V. E. Anderson, and R. H. Ritchie, Report No. RADC-TR-76-125 (1976).

[17] J. C. Ashley, C. J. Tung, R. H. Ritchie, and V. E. Anderson, Report No. RADC-TR-76-220, (June 1976).

[18] J. D. Jackson, Classical Electrodynamics, Chapter 13 2nd ed., John Wiley and Sons, Inc. New York, N. Y. (1975).

[19] Handbook for Auger Electron Spectroscopy, Physics Electronics Industries, Inc., (1976).

TABLE 3. AUGER MODEL FOR Au, Ag, Aℓ, AND C

MATERIAL	TRANSITION	ENERGY (keV)	YIELD
Au	M_1NN	2.111	0.35
		2.024	0.52
		1.772	0.13
	M_2NN	2.111	0.35
		2.024	0.52
		1.772	0.13
	M_3NN	2.111	0.35
		2.024	0.52
		1.772	0.13
	M_4NN	2.111	0.35
		2.024	0.52
		1.772	0.13
	M_5NN	2.111	0.35
		2.024	0.52
		1.772	0.13
	NNO	0.239	0.11
	NOO	0.069	0.89
Ag	LMM	2.570	0.90
	MNN	0.351	1.00
Aℓ	KLL	1.400	0.96
	LMM	0.070	1.00
C	KLL	0.272	1.00

region. The three pairs of lines arise from MNN transitions, the 0.239 keV line from an NNO transition, and finally the 0.069 keV line from an NOO transition. In our model given in Table 3, the pairs of lines have been combined giving us a total of five features. This is explained below.

The three MNN features arise from vacancies in the M_4 and M_5 subshells. There are no important lines associated with M_{123} NN transitions since vacancies in the M_1-M_3 subshells have a high probability of moving up to the M_4 and M_5 subshells. This is illustrated in Table 4 taken from McGuire.[20] ω_{M_i} is the probability of flourescence, a_{M_i} is the probability of a radiationless transition M_iXY with $X \neq M$ and $Y \neq M$, and $S_{M_{ij}}$ is the average number (not probability) of M_j vacancies occuring in a direct transition starting with an M_i vacancy. $S_{M_{ij}}$ will be greater than the corresponding probability of producing at least one M_j vacancy from an M_i vacancy when a super Coster-Kronig transition $M_iM_jM_j$ is energetically possible since, in number, this produces two M_j vacancies. Therefore, the sum

$$\omega_{M_i} + a_{M_i} + \sum_{j>i} S_{M_{ij}}$$

can be greater than unity as is the case for i=3 in the table.

Since a vacancy initially created in the M_1, M_2, and M_3 subshells transfers to the M_4 and M_5

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[20] E. J. McGuire, Phys. Rev. A5, 1043, (1972).

TABLE 4. TRANSITION PROBABILITY RELATED INFORMATION FOR THE M SHELL OF GOLD

(ω, a, and S are defined in the text.)

	ω_{M_i}	a_{M_i}	$S_{M_i,i+1}$	$S_{M_i,i+2}$	$S_{M_i,i+3}$	$S_{M_i,i+4}$
M_1	0.002	0.074	0.148	0.594	0.067	0.112
M_2	0.004	0.114	0.114	0.673	0.095	–
M_3	0.004	0.158	0.114	0.782	–	–
M_4	0.026	0.928	0.046	–	–	–
M_5	0.026	0.974	–	–	–	–

subshells with nearly unit probability, we can assign the $M_{45}NN$ Auger energies and yields to M_1, M_2 and M_3 vacancies. We have made this assignment in our model as illustrated in Table 3.

The final piece of input information is β, Burke's reflection coefficient. Values of β for the four materials appear in Table 5. We used Burke's empirical expression for β:

$$\beta = 0.475\ Z^{0.177} - 0.40 \qquad (22)$$

The basic quantity to be discussed in this section is the differential yield dY(E)/dE in units of electrons/photon-keV. For each of the four materials, we will provide comparisons of dY/dE for the source spectra shown in Figures 17-19. This will provide the first estimate of the sensitivity of photoemission to variations in the EWR source, as well as provide, in our opinion, the best available photoemission spectral information for Au, Ag, Aℓ, and C using this source.

We begin by considering the material Aℓ which has recently been studied in detail by us using our SXRP transport code under the DNA Contract DNA-001-77-0209. Since we have transport results available for Aℓ for spectra 2, 3 and 4 given in Figures 18 and 19, we first present these in Figure 20. All photoemission spectra are dominated by KLL Auger electrons whose initial

TABLE 5. BURKE'S EMPIRICAL
PARAMETER β VERSUS Z

Z	MATERIAL	β
6	Carbon	0.25
13	Aluminum	0.35
47	Silver	0.54
79	Gold	0.63

energy is 1.4 keV. The low energy peak is due to LMM Auger electrons and K-photoelectrons. Empirical results for all four source spectra are shown in Figure 21. A significant decrease in the photoemission yield is observed for spectrum 4 because this spectrum contains ∿50% of its photons below the Aℓ 1.56 keV K-edge. Such photons produce little photoemission.

By including the transport results shown in Figure 20, we are able to demonstrate the degree of accuracy of the empirical spectra in Figure 21. We note that there is qualitative agreement, but significantly less structure in the empirical results. The greater structure in the transport results is caused by elastic scattering which is not treated by the empirical method. A detailed discussion of this effect has been given in Section 3. The following word of caution should be given — the degree of difference between empirical and transport results can be expected to be material dependent. We thus reserve comment on the accuracy of the empirical photoemission spectra for the other materials which have not yet been modeled for transport calculations.

The next material we consider is gold. Because it is a complicated material to model, our first empirical results were obtained for monoenergetic photon sources for the purpose of comparison with available data. This comparison is given in Figure 22 in terms of total yield in electrons/photon from $h\nu$ = 10 keV down to 0.1 keV. The data were taken from Figure 15 of Burke's paper.[4]

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[4] E. A. Burke, "Soft X-Ray Induced Electrom Emission," IEEE Trans. Nuc. Sci., NS-24, No. 6, 2505 (1977).

In general, our empirical results are ∿25% below both the data and Burke's results. We have not resolved the discrepancy as of this time. Part of the difference will be due to more detail given to the subshell contributions in our formulation.

The yield dY/dE for the four source spectra for Au obtained by the empirical code appear in Figure 23. In general, the spectra are harder for Au than for Aℓ and show more sensitivity to the source spectra. The softer, less structured yields for photon spectra 3 and 4 result from these spectra having no photons above 2.2 keV and thus not allowing for M shell ionization. To provide insight into the results for Au, dY/dE for source spectrum 2 is shown in Figure 24 in terms of its many components for Auger and photoelectrons. We see that the Au spectrum is harder than the Aℓ spectrum because of the MNN Auger and N shell photoelectron emission.

The empirical results for Ag and C appear in Figures 25 and 26. We observe that the emission spectra are not very sensitive to differences between the first three source spectra, except at high electron energies.

To conclude this section, we give the integrated yields in Table 6 for the four materials and four source spectra. Au and C provide the highest and lowest yields, respectively. The Au yields may be too low based on the 25% difference between the empirical results and the data for line sources shown previously in Figure 22. For Aℓ the empirical yield is ∿30% higher than the transport yield. In general, the numbers in this table should be

TABLE 6. TOTAL YIELDS IN ELECTRONS/PHOTON
FOR THE SOURCE SPECTRA IN FIGURES 1-3

Multiply the numbers by 10^{-3}.

	SPECTRUM			
	1	2	3	4
Aℓ-TRANSPORT	–	7.0	7.1	5.6
Aℓ-EMPIRICAL	9.9	10.1	10.1	6.2
Au-EMPIRICAL	13.2	12.5	8.9	11.3
Ag-EMPIRICAL	7.4	7.5	7.6	9.8
C-EMPIRICAL	0.68	0.69	0.71	0.90

accurate to 50% or less based on Burke's success in reproducing total yield data for a wide variety of materials.

One must not draw general conclusions from Table 6 about relative emission efficiencies from one material to the next. One such mistaken conclusion might be that aluminum emits photoelectrons nearly as efficiently as gold. In general, gold is a better photoemitter. Aluminum happens to be a strong photoemitter for EWR sources since a significant fraction of the photons lies just above the Aℓ K-edge. If the EWR spectrum were shifted as little as 0.2 keV toward lower energy, the Aℓ yield would drop more than a factor of two. It should be emphasized that the photoemission efficiency depends critically on the distribution of photons with respect to the photoabsorption edges for a given material.

To summarize this section, an empirical photoemission code has been developed to provide an additional means of making comparisons with SXRP results and to provide estimates of photoemission yields for materials not yet treated by the SXRP code. The purpose of this section has been to document the contents of the empirical code and present our first analysis using this code. The needed atomic information has been presented for the four materials considered — Au, Ag, Aℓ, and C. For Au, Ag, and C, the given data represents our first attempt to model these materials and, consequently, may undergo revisions in anticipated future analyses.

Differential yields were presented for four different source spectra to establish the sensitivity of the photoemission spectrum, as well as total yield to variations in the EWR source. Empirical results were presented for all four materials. Transport results were also presented for Aℓ, not only to show sensitivity to the source spectrum, but also to compare with the empirical emission spectra.

Gold was shown to possess the hardest spectrum and highest yield. It also exhibited a greater sensitivity to the various source spectra considered than the other materials. We believe the transport results for Aℓ and empirical results for Au, Ag, and C presented in this report to be the best such information currently available for an EWR source.

Section 6

CHARGE DEPOSITION PROFILES NEAR ^{60}Co IRRADIATED MATERIAL INTERFACES

Recent experimental work by Frederickson[6] has provided the users of transport codes with a large data base which may be used for code validation of charge deposition profiles near irradiated material interfaces. Results were presented for charge deposition in thin foils (<0.004 electron range) of Aℓ, Sn, or Ta between various combinations of equilibrium thick slabs of Aℓ, Cu, Sn, or Pb. The radiation source was ^{60}Co γ-radiation.

Calculations of a few configurations have been performed using the one-dimensional version of the POEM[25] code. A diagram of the transport geometry is shown in Figure 27. For these calculations, the photons were assumed normally incident upon the surface. The energy spectrum used was the test spectrum integrated over angle.[6] The configurations considered are shown in Table 7.

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[6] A. R. Frederickson, "Charge Deposition Profiles Near Irradiated Material Interfaces," IEEE Trans. Nuc. Sci., NS-23, No. 6, 1867 (1976).

[25] W. L. Chadsey, "POEM: A Fast Monte Carlo Code for the Calculation of X-Ray Photoemission and Transition Zone Dose and Current," AFCRL-TR-75-0323 (15 January 1975).

TABLE 7. CONFIGURATIONS FOR CALCULATIONS

SLAB 1	FOIL	SLAB 2
Aℓ	Aℓ	Pb
Pb	Aℓ	Aℓ
Pb	Sn	Sn
Sn	Sn	Aℓ
Aℓ	Ta	Aℓ
Sn	Ta	Pb

The charge deposition calculations for the given configurations provide a very stringent test of the transport code. The charge deposition is given as a difference between two values of the current which are nearly equal as a result of the thinness of the foils. Agreement between the calculations and measurements, thus provides a high degree of confidence in the POEM code.

For each configuration, six components of the current were calculated as illustrated in Table 8 for the Aℓ/Aℓ/Pb configuration. The net current in the foils is given by

$$J_{net} = J_{eq}+(J_1-J_2)+(J_3-J_4)+(J_5-J_6) \qquad (23)$$

where J_{eq} is the equilibrium current in Aℓ and the subscripts refer to the components defined in Table 8. The charge deposition in each foil was obtained as the difference in current at the left and right faces of the foil. In general, the net current J_{net} is determined by small differences between the dominant components which are of nearly the same magnitude. Furthermore, the charge deposition in a given foil is the difference in net current at the two surfaces of the foil. The difference is small since the foils are thin. Therefore, high accuracy is required in the Monte Carlo calculations of the current. The current contributions 1 through 6 are shown in Figure 28 for the Aℓ/Aℓ/Pb configuration.

In order to accurately calculate the charge deposition, a minimum of 20,000 electron histories were

TABLE 8. COMPONENTS OF CURRENT

COMPONENT	SLAB 1	FOILS	SLAB 2	ELECTRON SOURCE
1	Aℓ	Aℓ	Pb	Emission from Slab 2 (Pb)
2	Aℓ	Aℓ	Aℓ	Emission from Slab 2 (Aℓ)
3	Aℓ	Aℓ	Pb	Backscatter from Slab 2 (Pb) of electrons emitted from foils
4	Aℓ	Aℓ	Aℓ	Backscatter from Slab 2 (Aℓ) of electrons emitted from foils
5	Aℓ	Aℓ	Pb	Emission from Slab 1 (Aℓ)
6	Aℓ	Aℓ	Aℓ	Emission from Slab 1 (Aℓ)

followed for each contribution. Each of these contributions was fitted within a function of the form

$$J = J_o \exp(Ax+Bx^2+Cx^3) \qquad (24)$$

and the coefficients were determined by a least squares fit to the calculated results. The fits were then evaluated to determine the charge deposition.

The results for the first five configurations shown in Table 7 are presented in Figures 29-33 with the same units used in Reference 6. The histograms are the POEM results and the data points with error bars are the Frederickson results. The agreement between the experimental results and the calculations is, in general, quite good; the results of the calculation in most cases fall within the experimental error bars. Where discrepancies occur, they appear to be statistical in nature, such that the inclusion of more histories in the Monte Carlo calculation would reduce the discrepancies. The calculational results for the Sn/Ta/Pb configuration were quite noisy, in some instances having a different sign than the data. The charge deposition is small in this case, but about the same magnitude found for the Pb/Aℓ/Aℓ configuration. The reason for this discrepancy is not known at this time.

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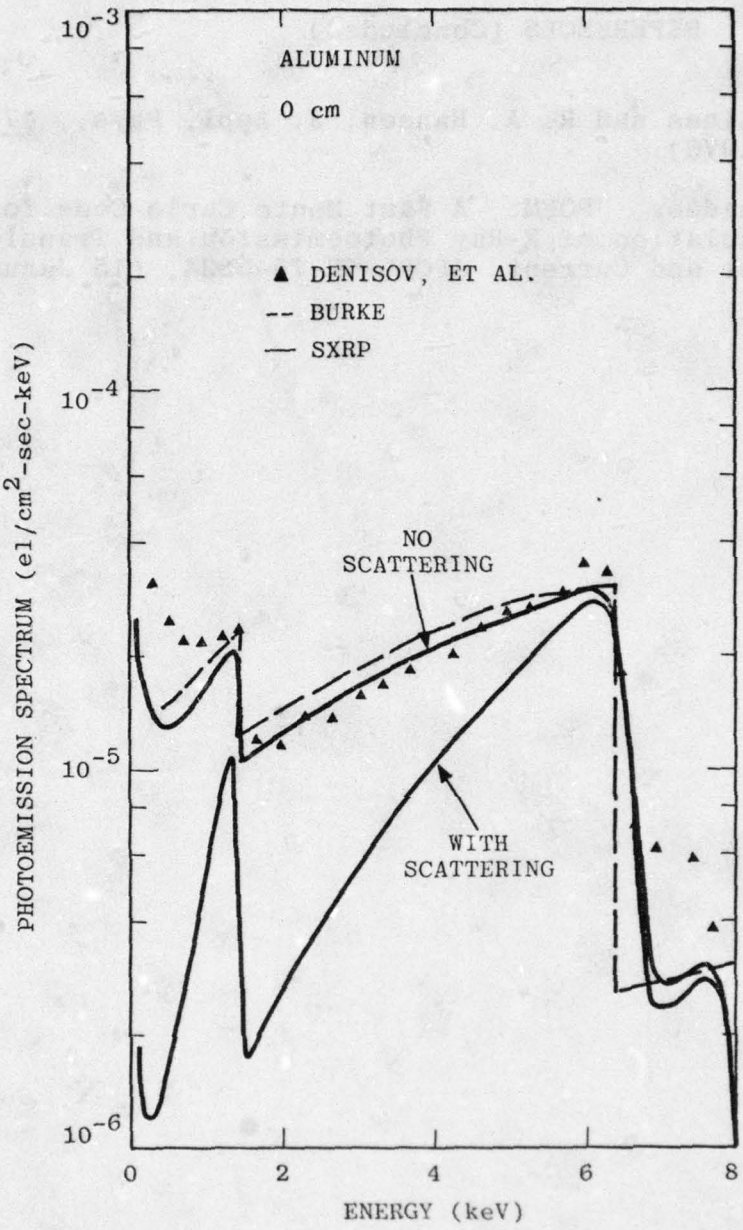

FIGURE 1. Photoemission Spectra for an 8 keV Photon Source Incident on Aℓ. The Solid Curves are SXRP Results With and Without Scattering

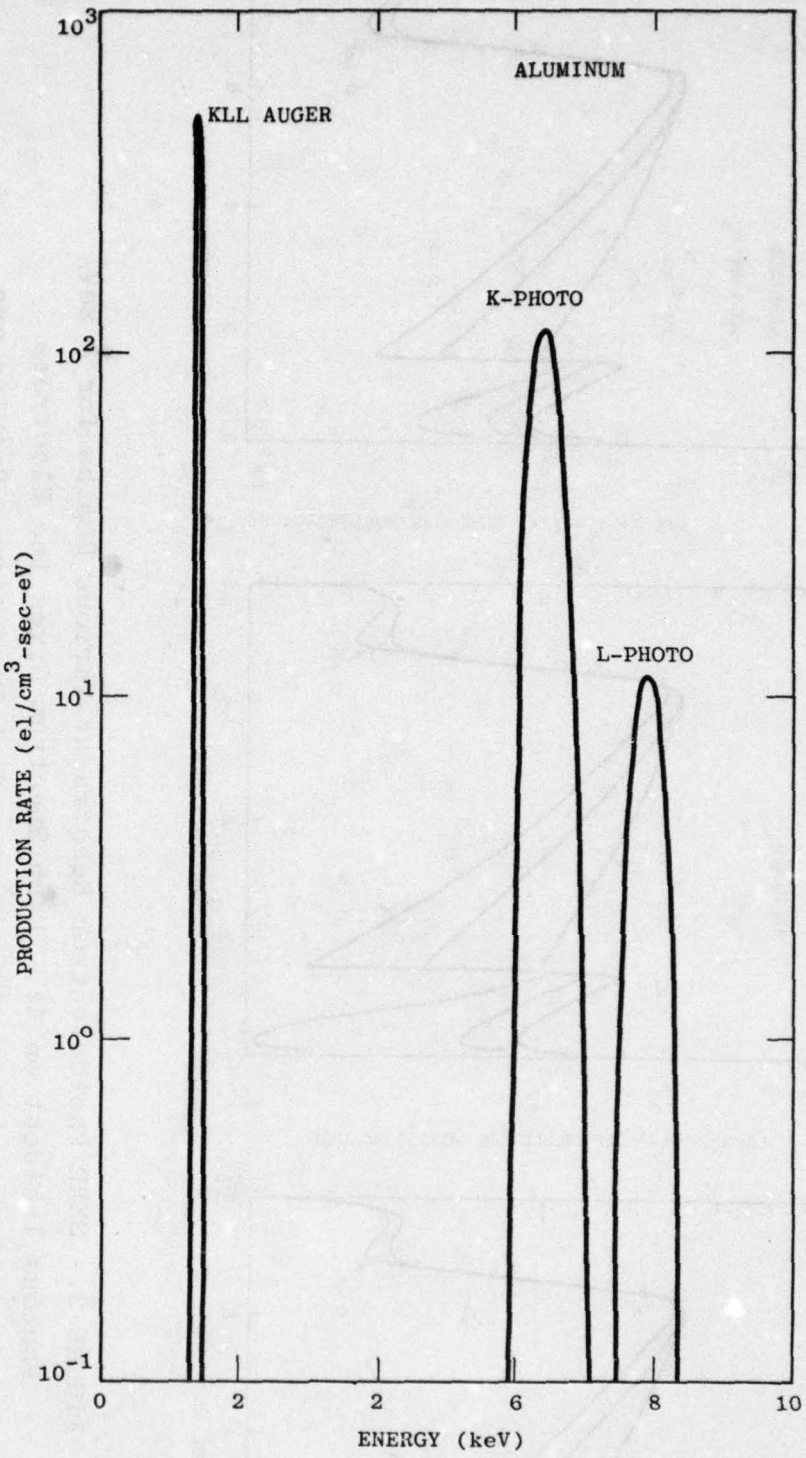

FIGURE 2. Source Distribution of Photo- and Auger Electrons for the 8 keV Gaussian Photon Source Used to Perform the SXRP Calculations

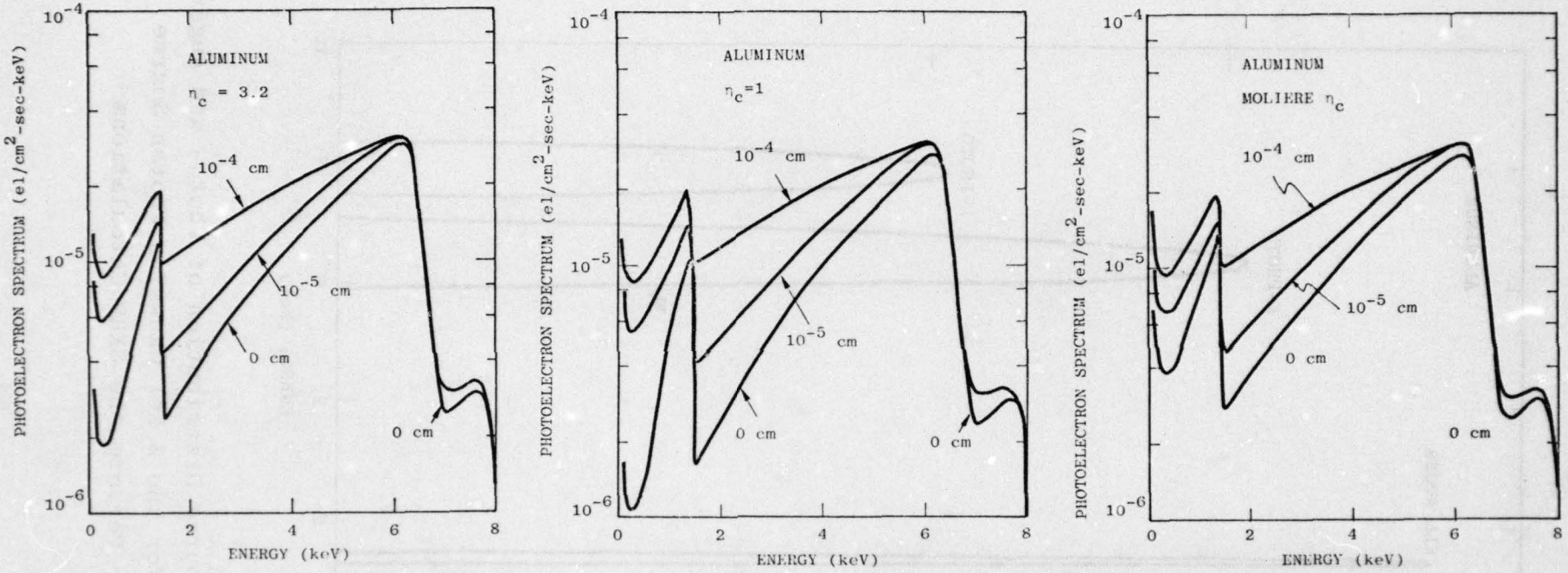

FIGURE 3. SXRP Photoelectron Spectra at Various Depths for 8 keV Photons Incident on Aℓ. Each Spectrum Gives the Electrons Passing Through a Unit Area in the Back Direction. Results are Shown for Three Different Representations of the Elastic Scattering IMFP Designated by η_c

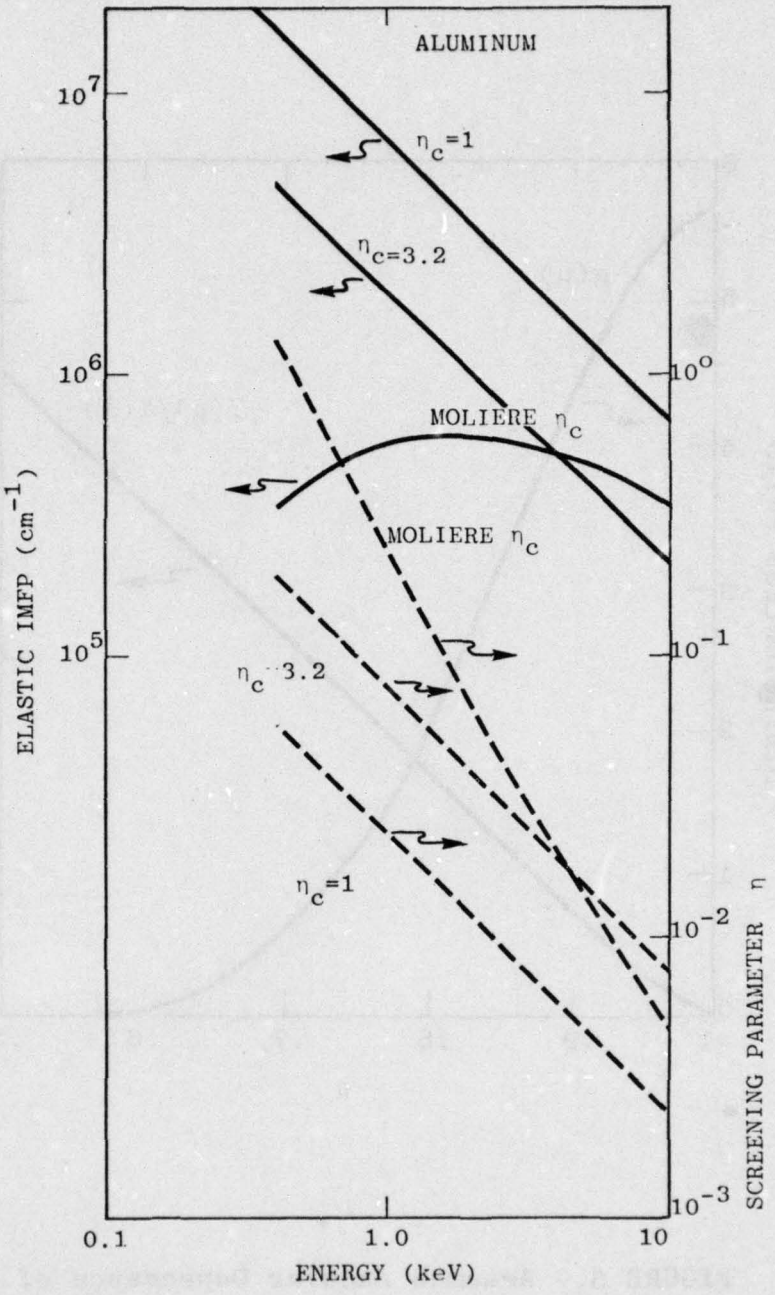

FIGURE 4. Energy Dependence of Screening Parameters and Corresponding Elastic IMFP's Used to Perform the Backscatter Calculations for Aℓ

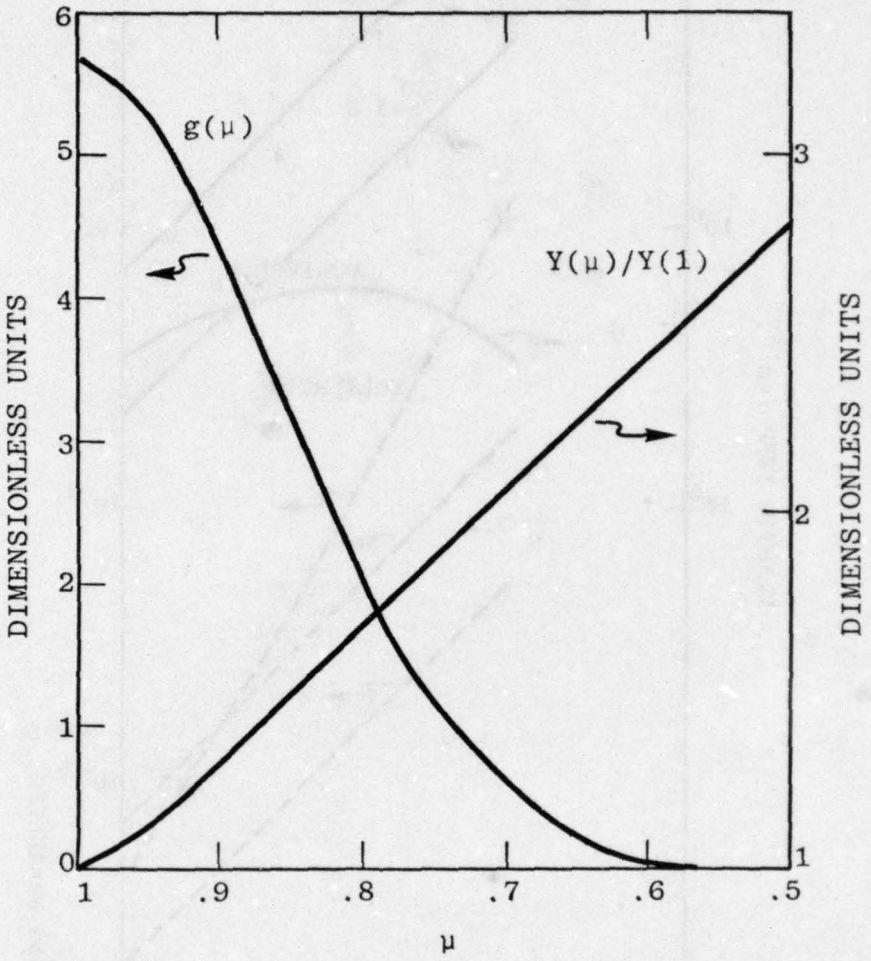

FIGURE 5. Assumed Angular Dependence of the Incident Electrons and the Backscatter Yield. μ is the Cosine of the Incidence Angle

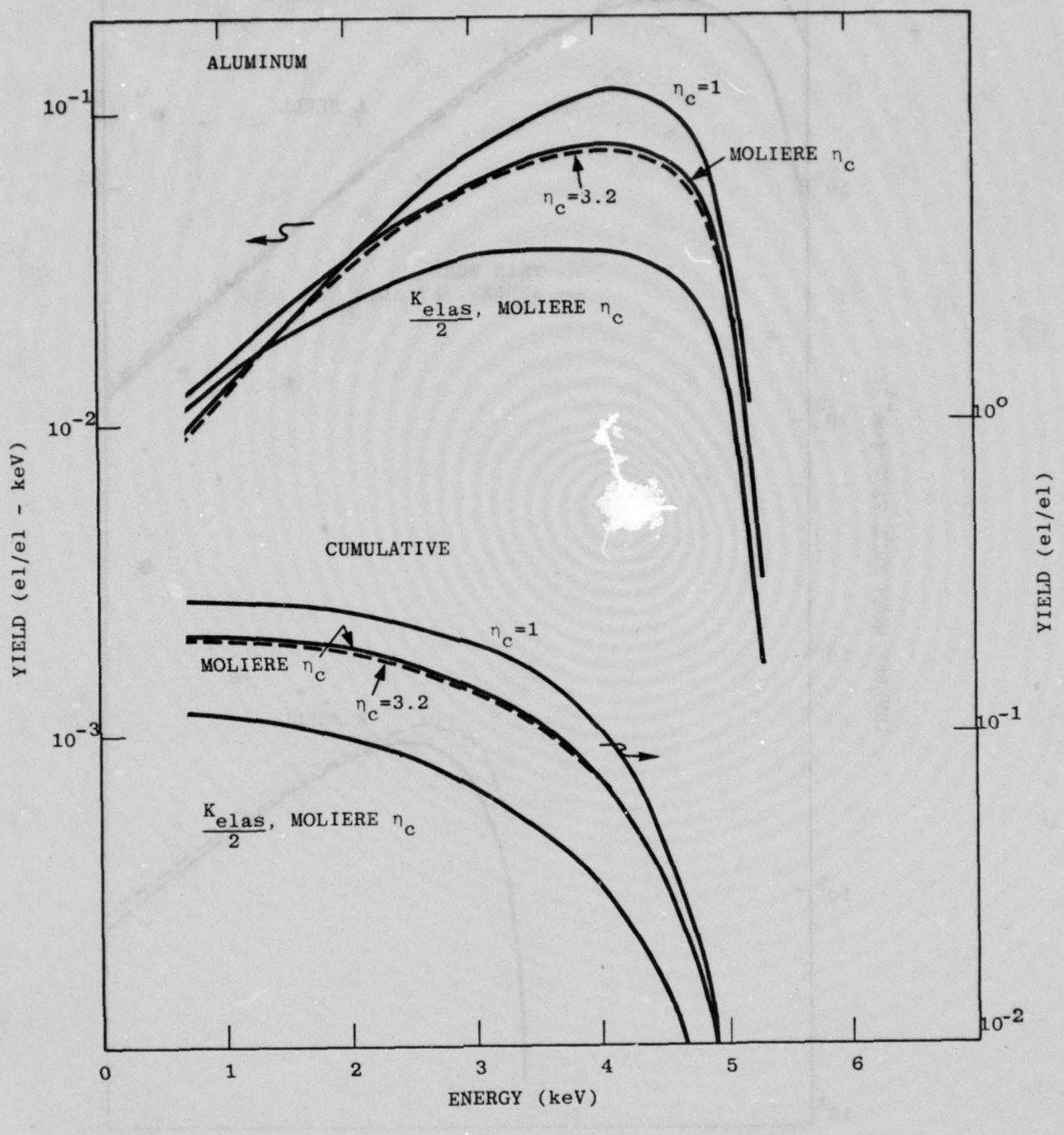

FIGURE 6. Differential and Cumulative Backscatter Yields for the Various Representations of the Elastic Scattering Cross Sections Described in the Text

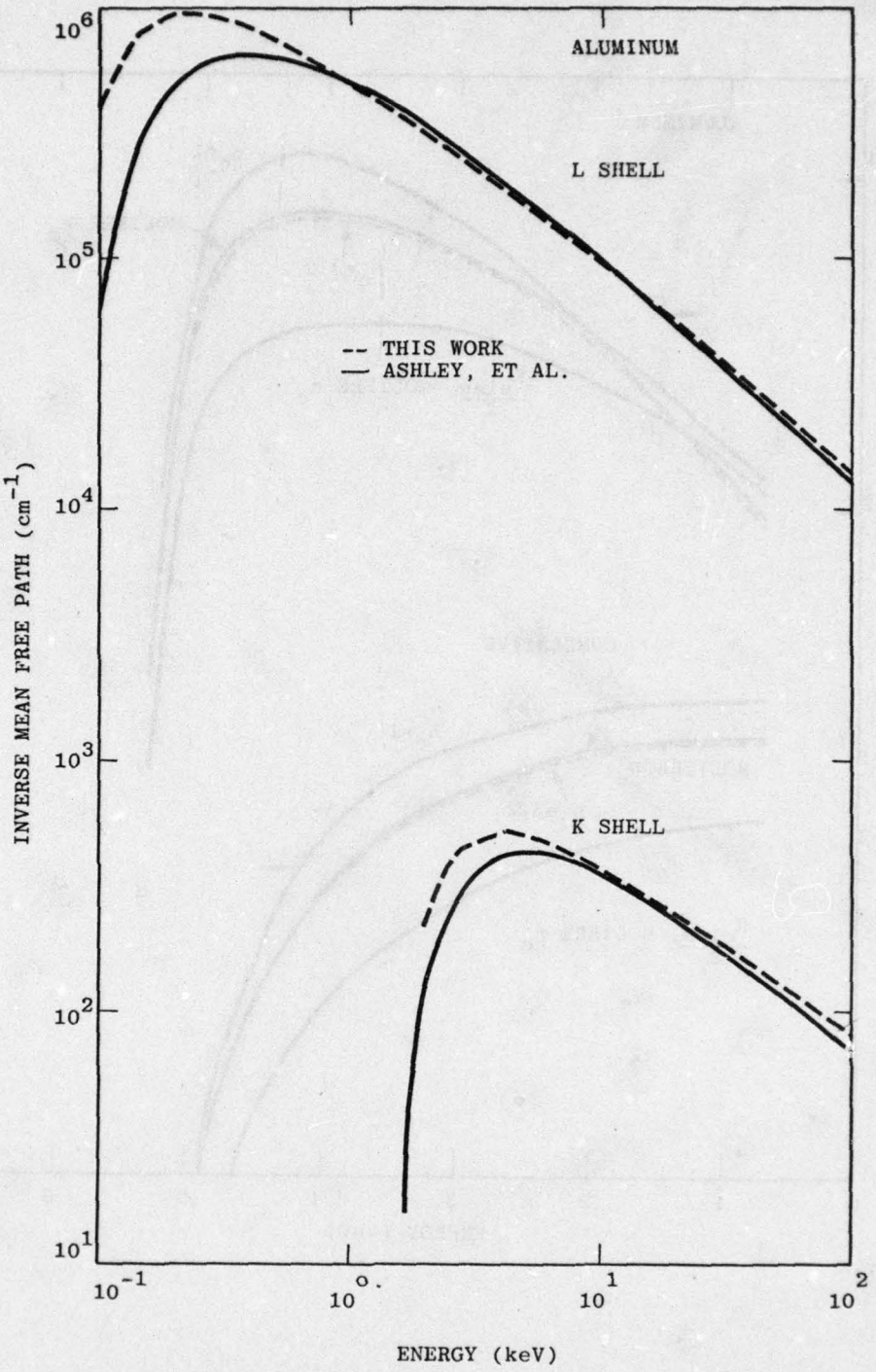

FIGURE 7. A Comparison of Inner Shell IMFP's Given by Integrating Equation (10) and by Ashley, et al.[11] from more rigorous calculations

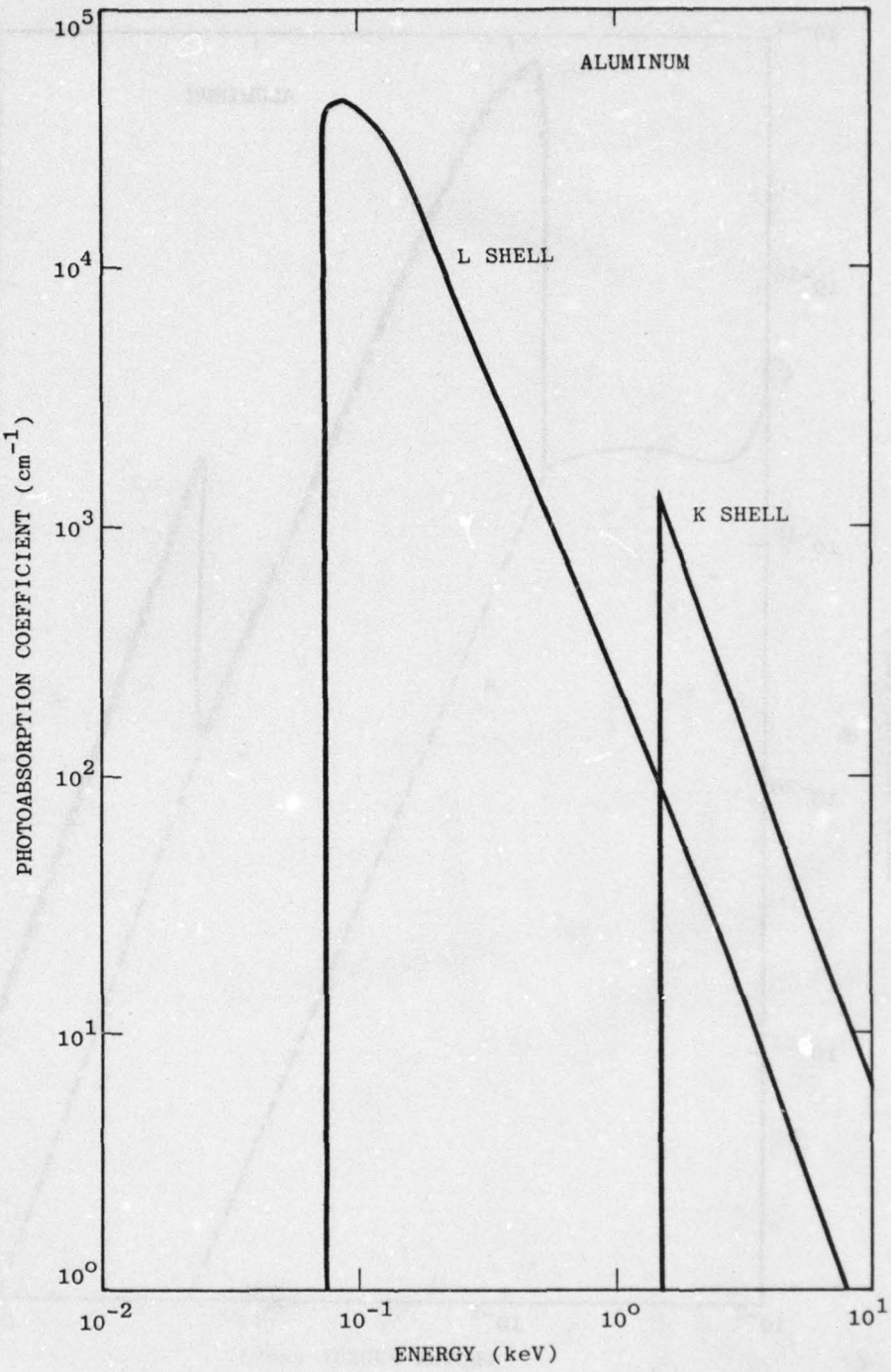

FIGURE 8. The Photoabsorption Coefficients Used to Calculate the IMFP's shown in Figure 7

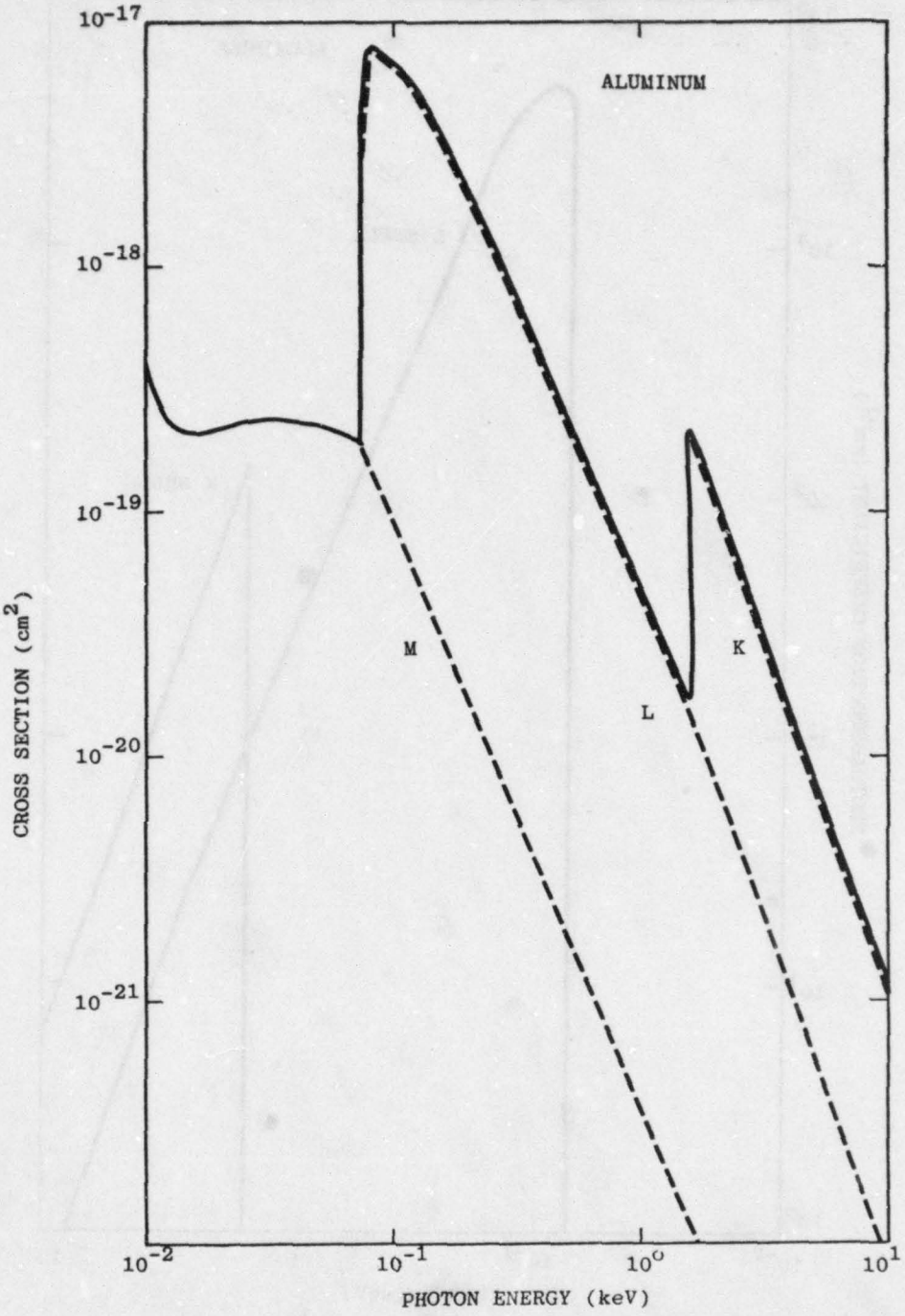

FIGURE 9. Photoabsorption Cross Sections for Aℓ

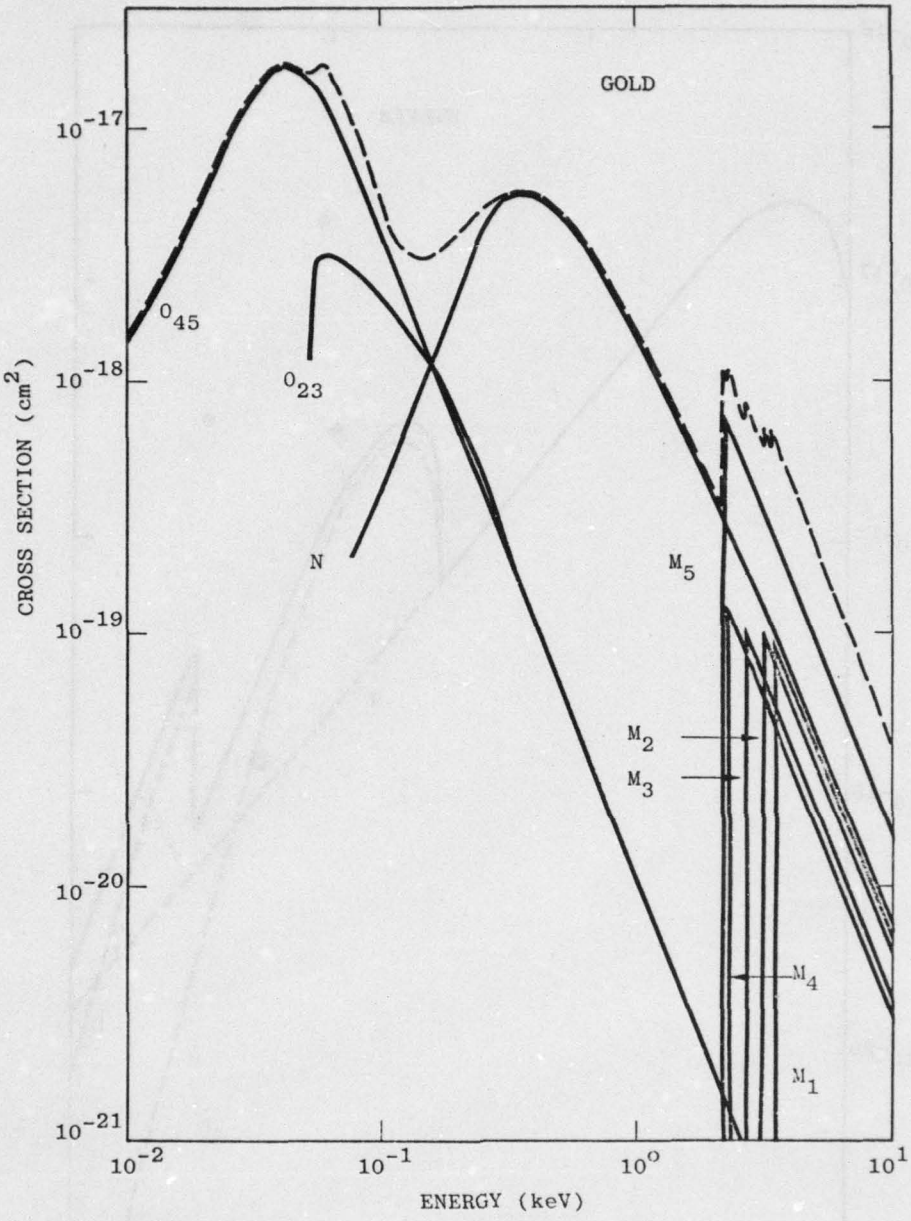

FIGURE 10. Photoabsorption Cross Sections for Au

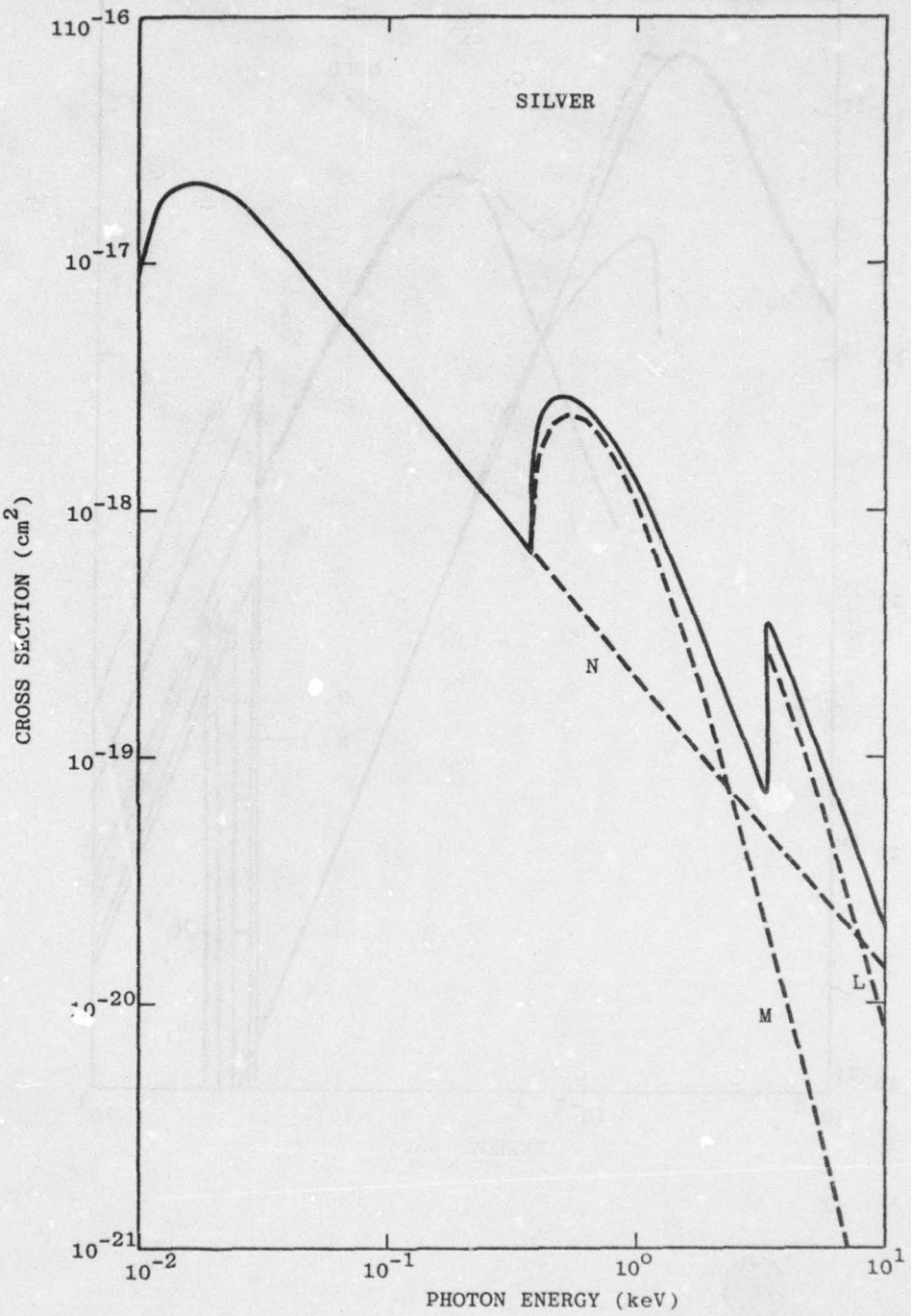

FIGURE 11. Photoabsorption Cross Sections for Ag

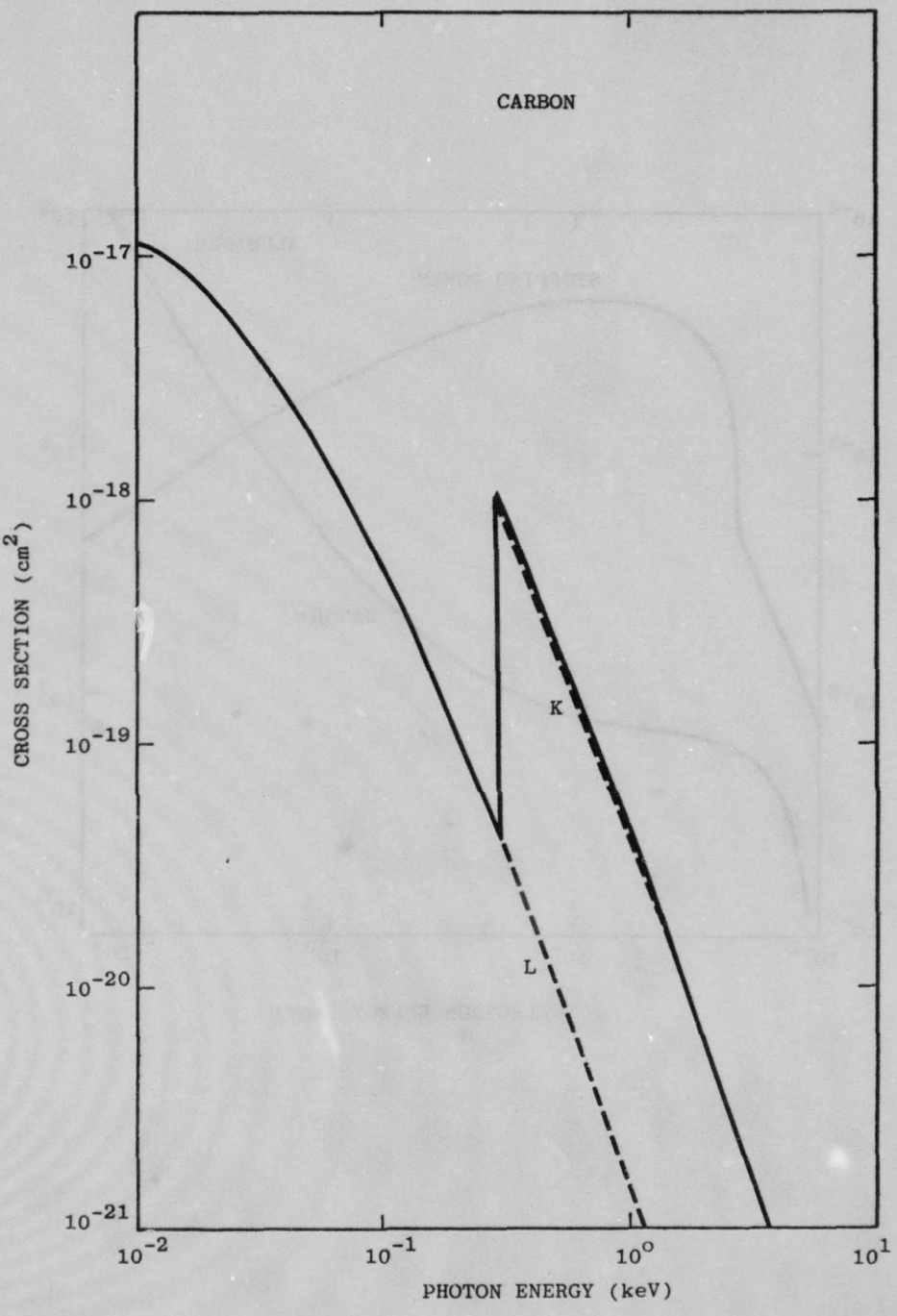

FIGURE 12. Photoabsorption Cross Sections for C

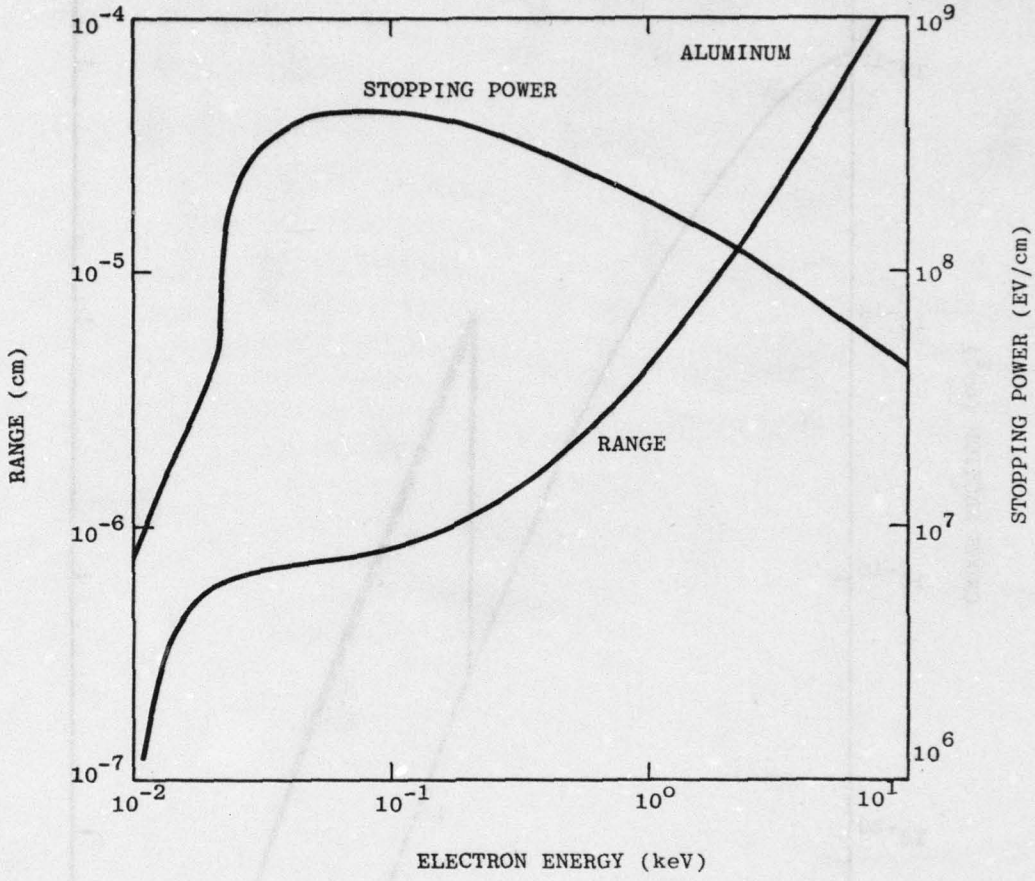

FIGURE 13. Range and Stopping Power for Aℓ

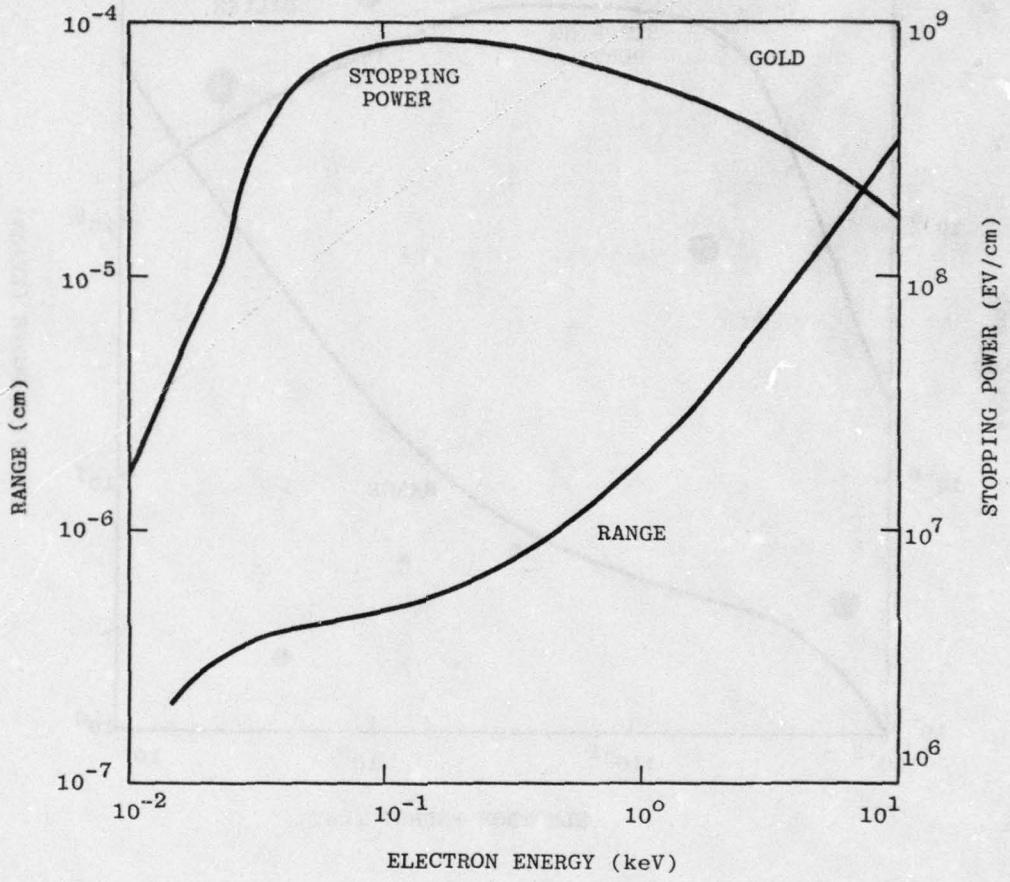

FIGURE 14. Range and Stopping Power for Au

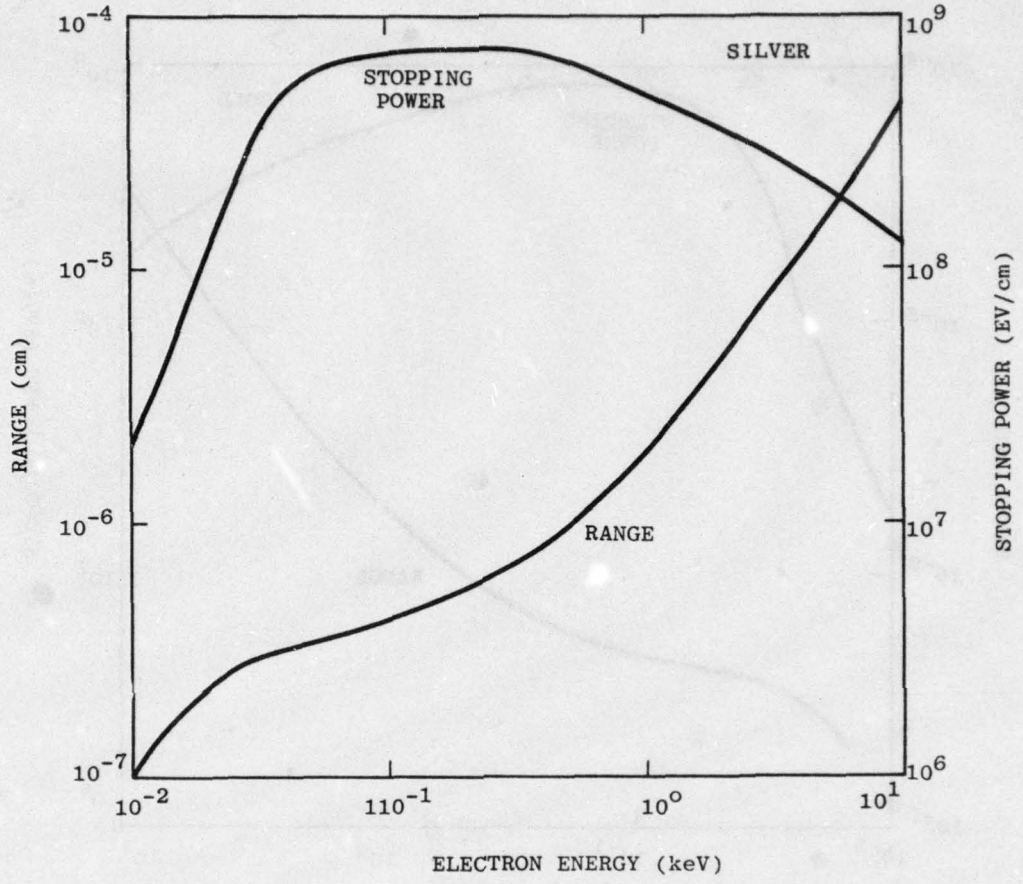

FIGURE 15. Range and Stopping Power for Ag

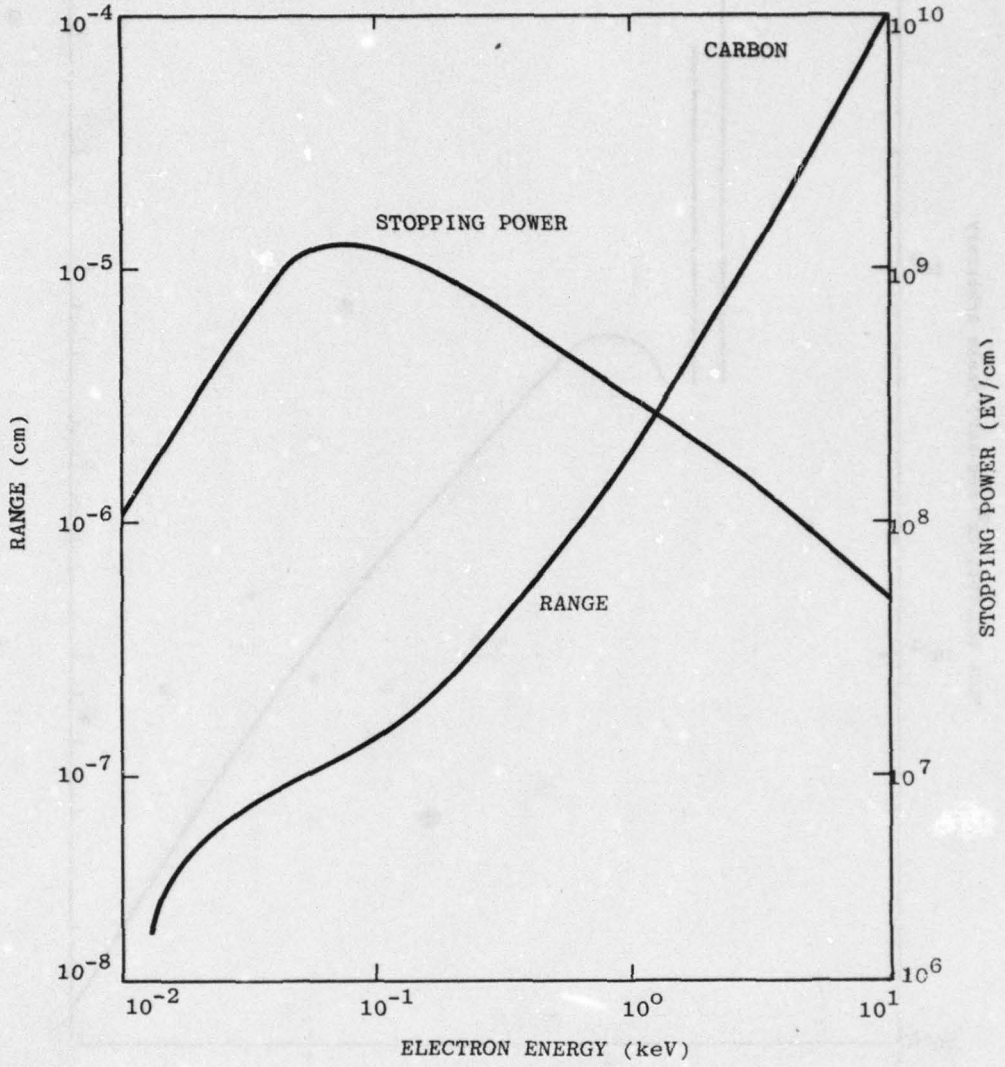

FIGURE 16. Range and Stopping Power for C

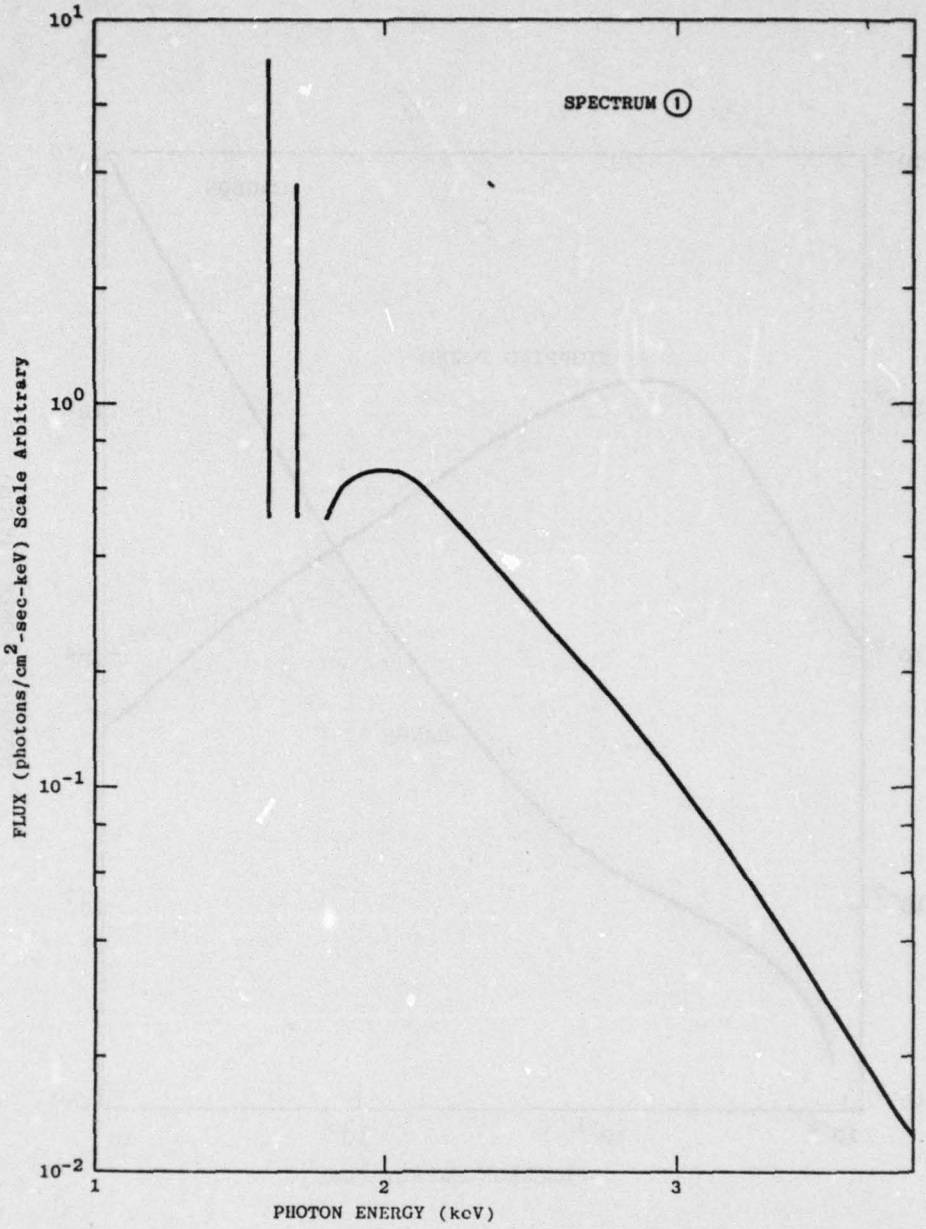

FIGURE 17. The First of Four Photon Spectra Used to Study the Sensitivity of Photoemission to Variations in the EWR Source

Approximately 50% of the photon energy is contained in the two lines.

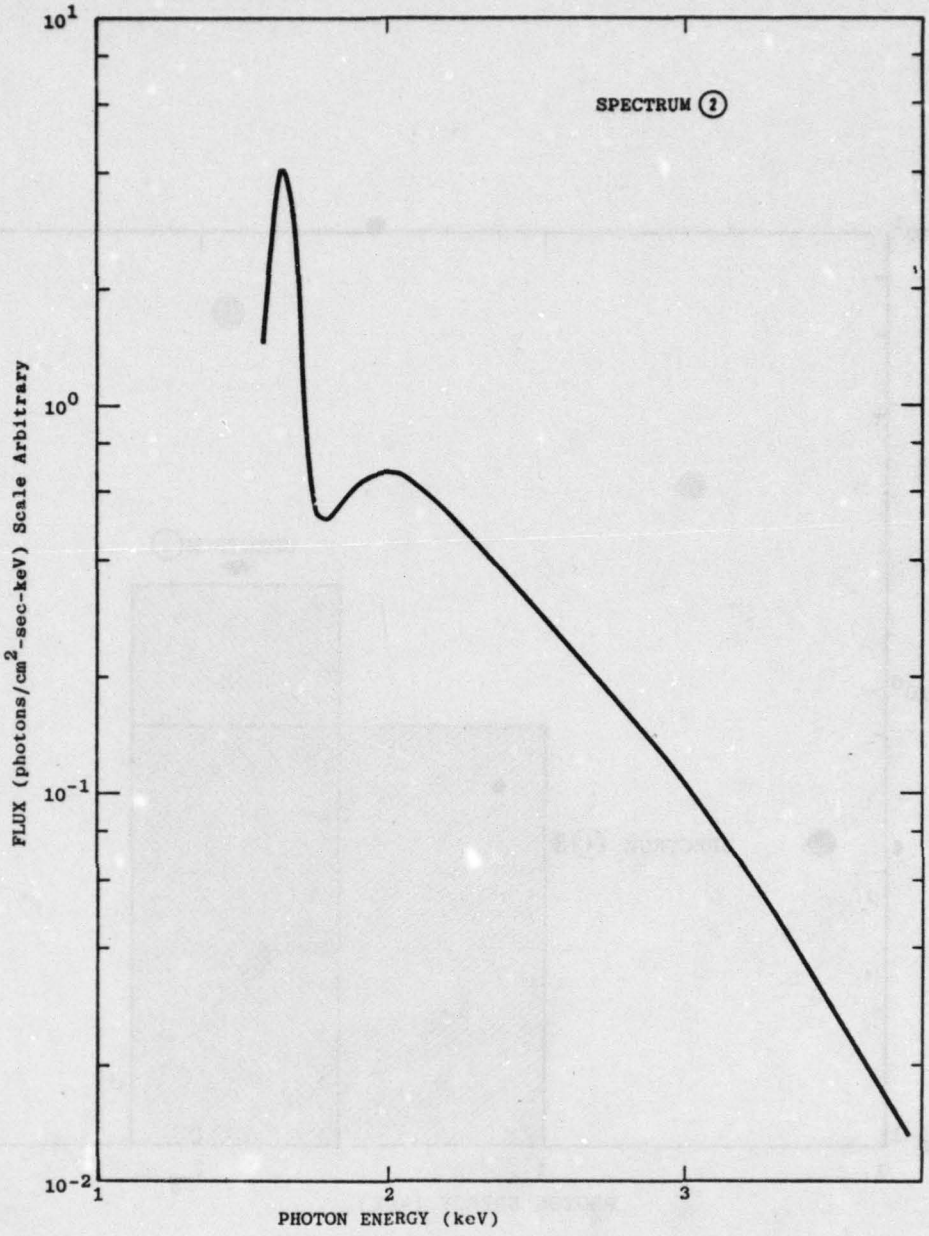

FIGURE 18. The Second of Four Photon Spectra Considered

The Gaussian feature centered at 1.6 keV contains ∿50% of the photon energy and is used as a continuum approximation for the two lines in spectrum (1).

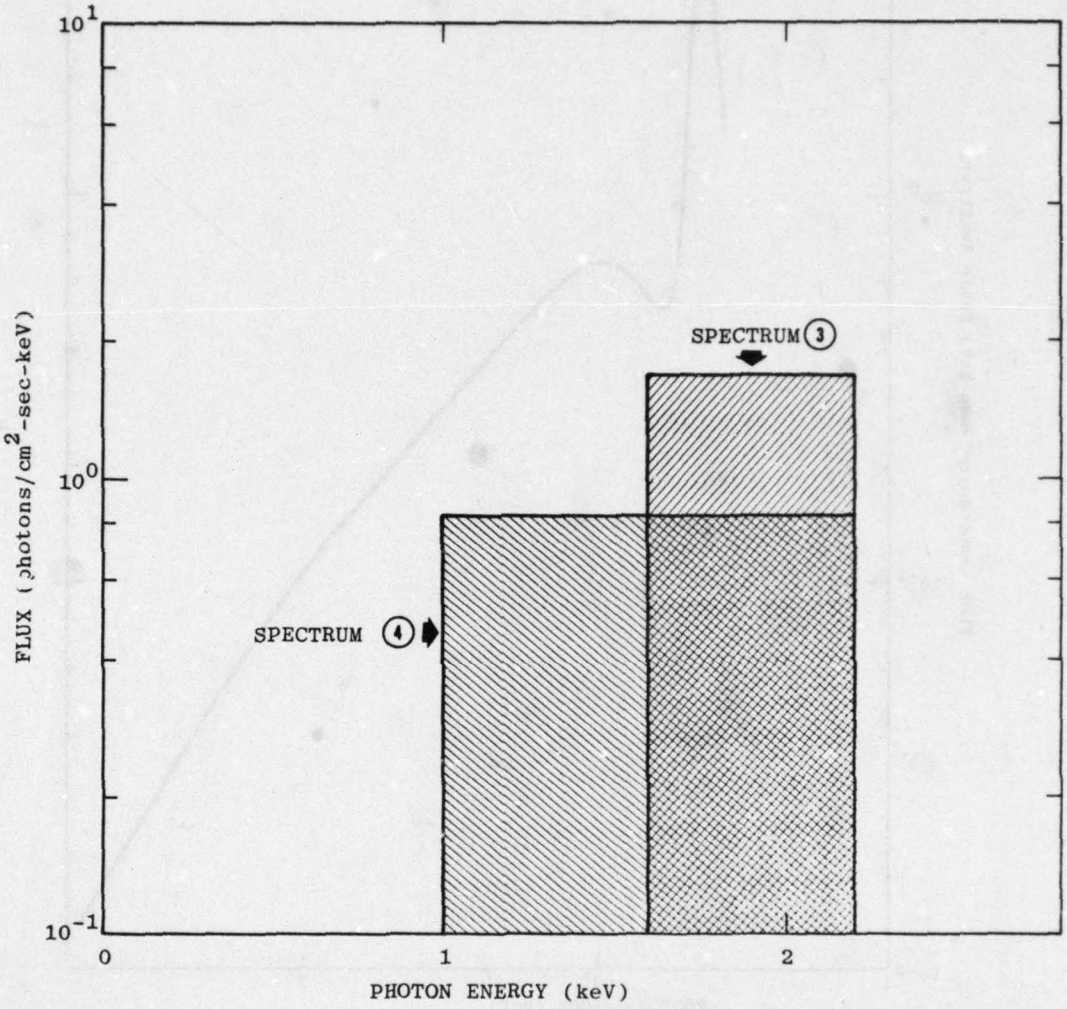

FIGURE 19. The Remaining Two Photon Spectra for Performing the Sensitivity Analysis

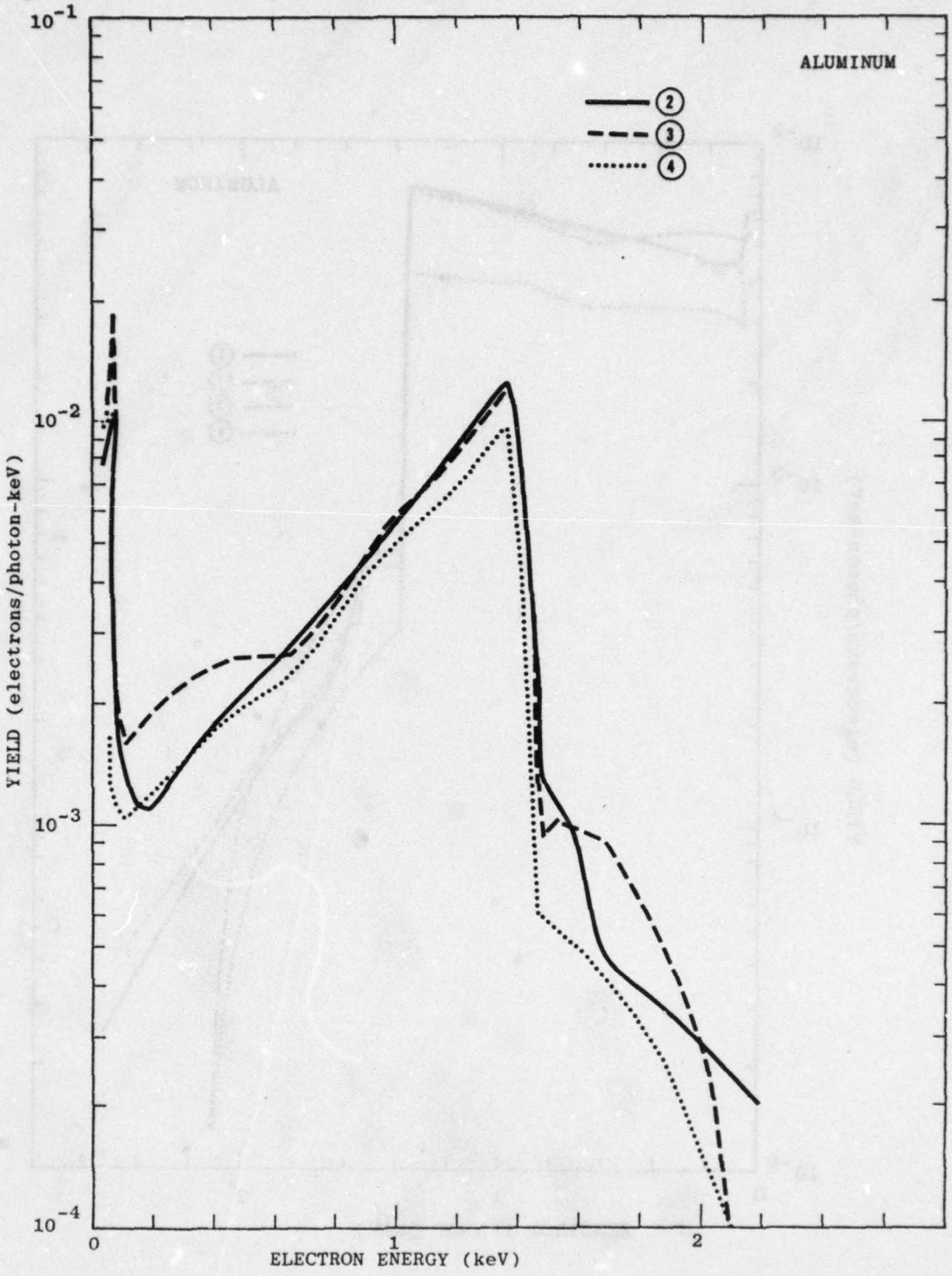

FIGURE 20. Differential Yields for Aℓ Obtained with the Transport Code for Photon Spectra (2)-(4)

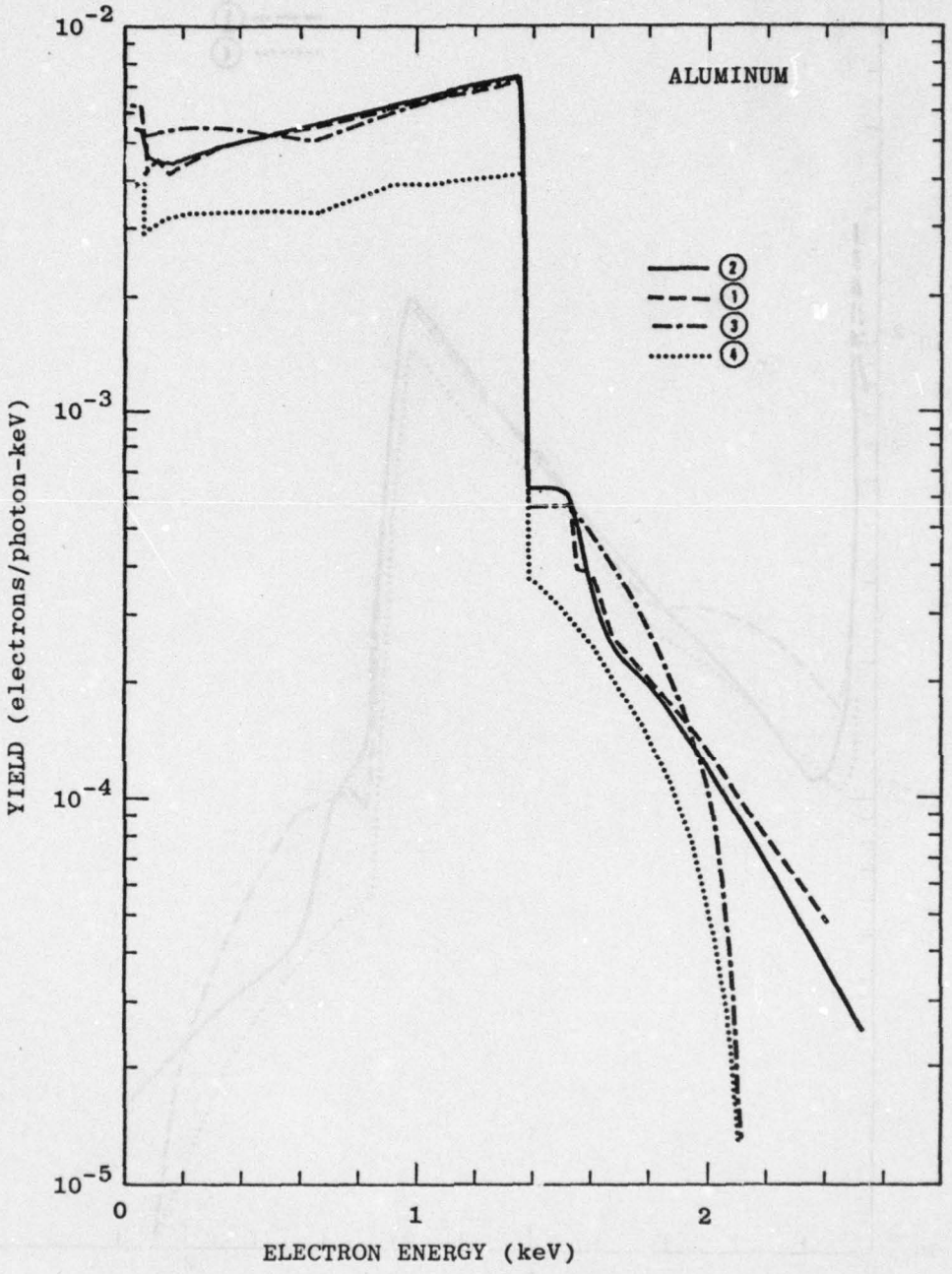

FIGURE 21. Empirical Differential Yields for Aℓ for Photon Spectra ①-④

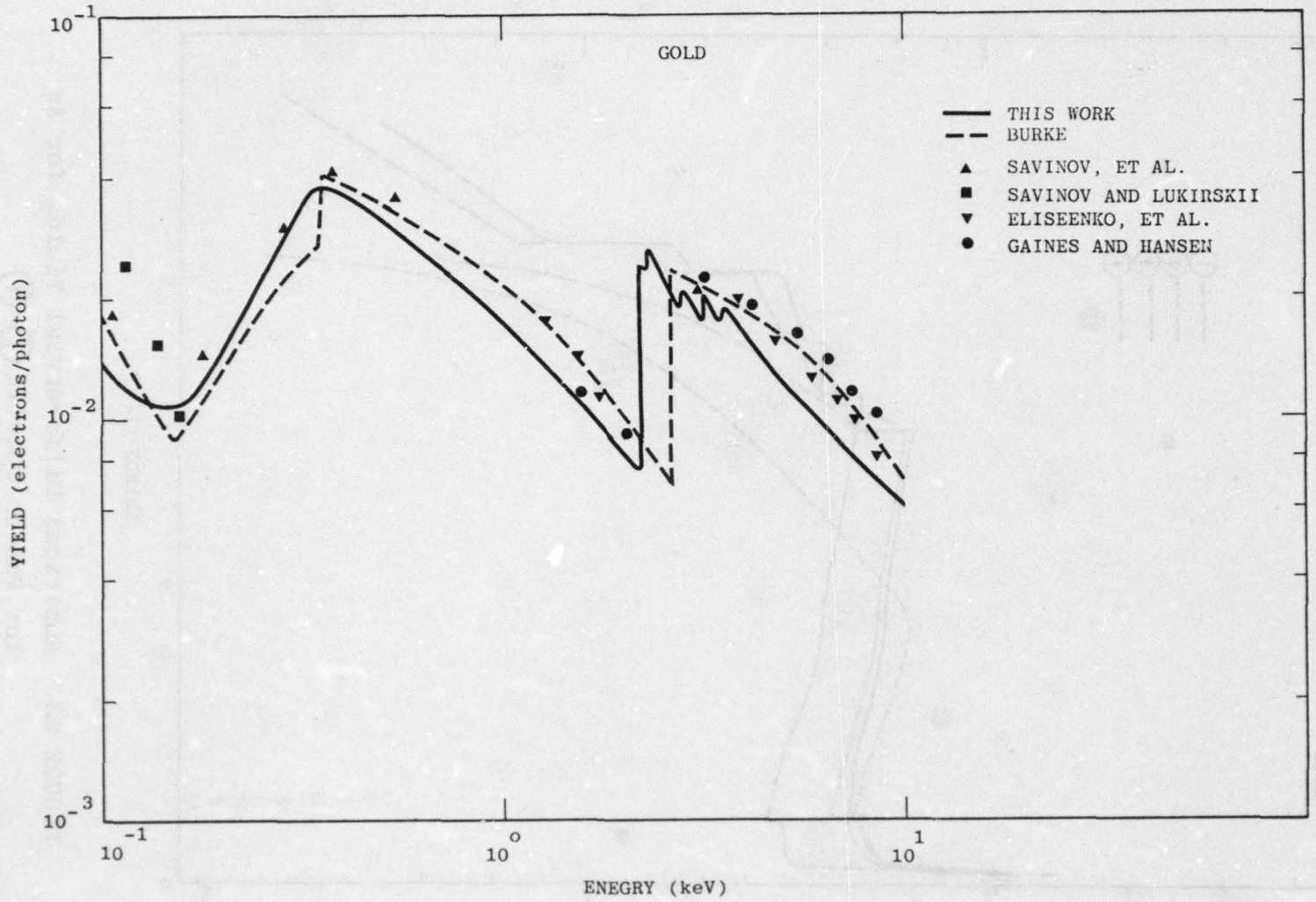

FIGURE 22. Empirical Total Yields versus Photon Energy for Au

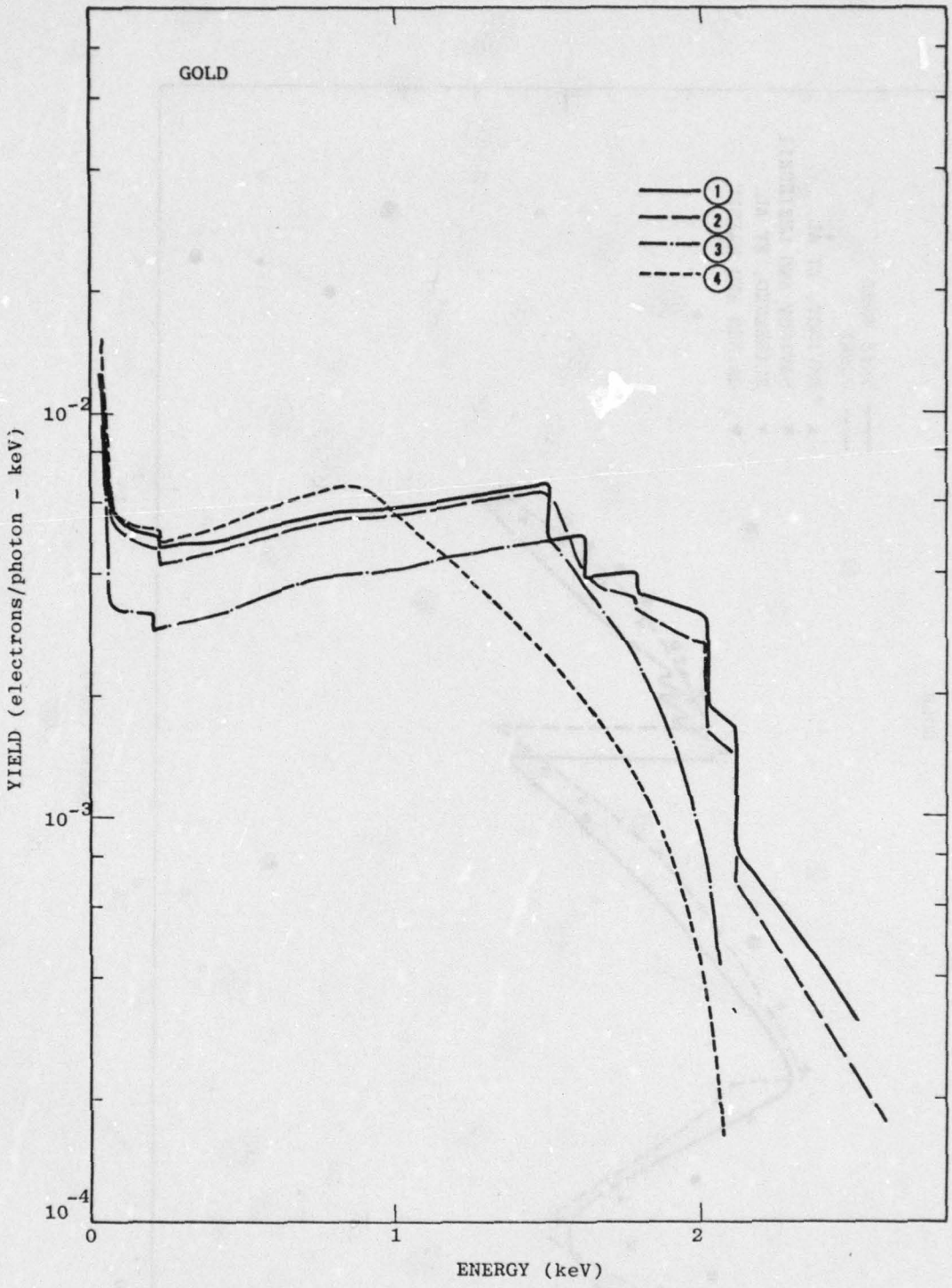

FIGURE 23. Empirical Differential Yields for Au for Photon Spectra (1)-(4)

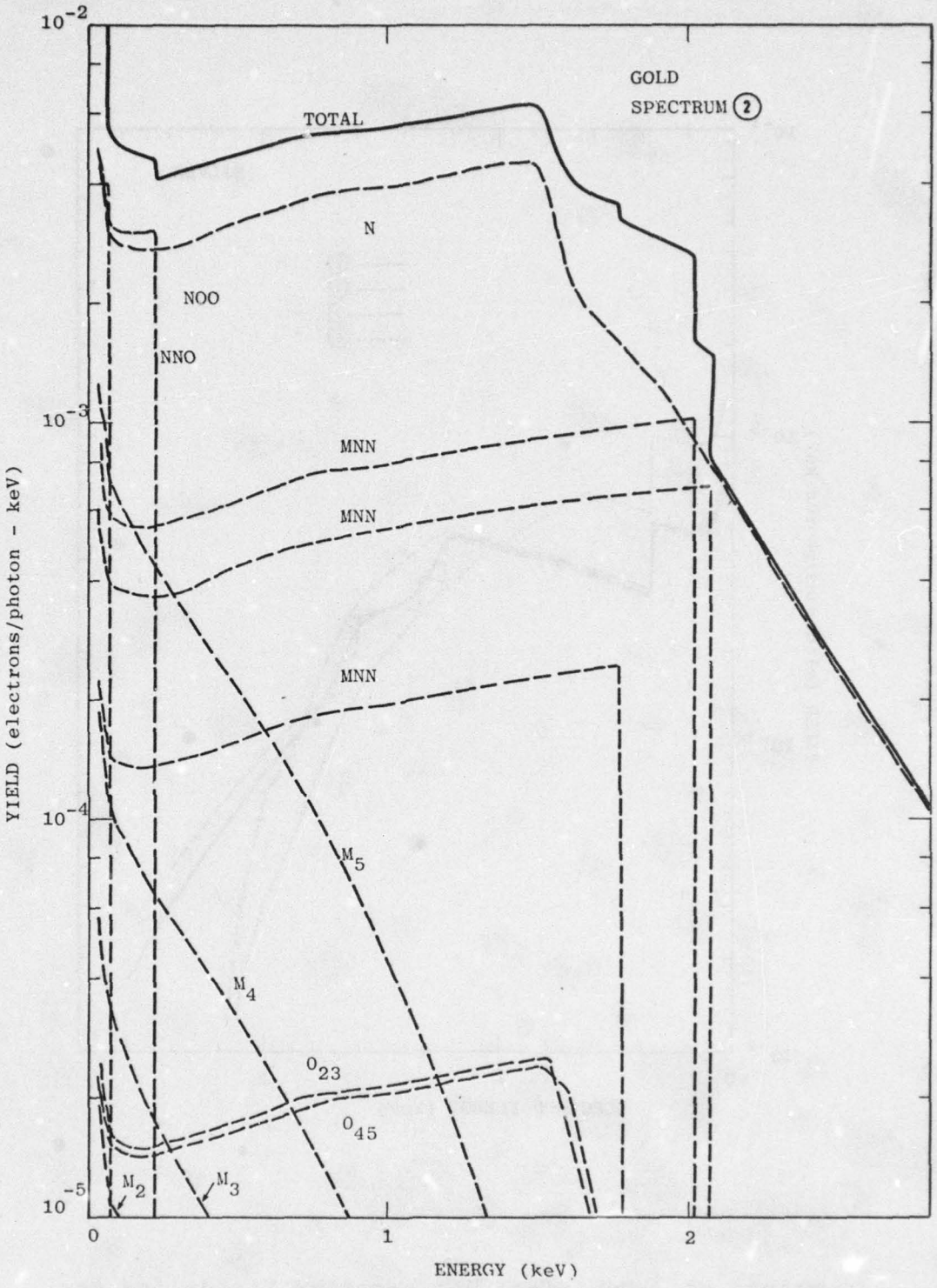

FIGURE 24. Empirical Differential Yield Components for Au for Photon Spectrum (2)

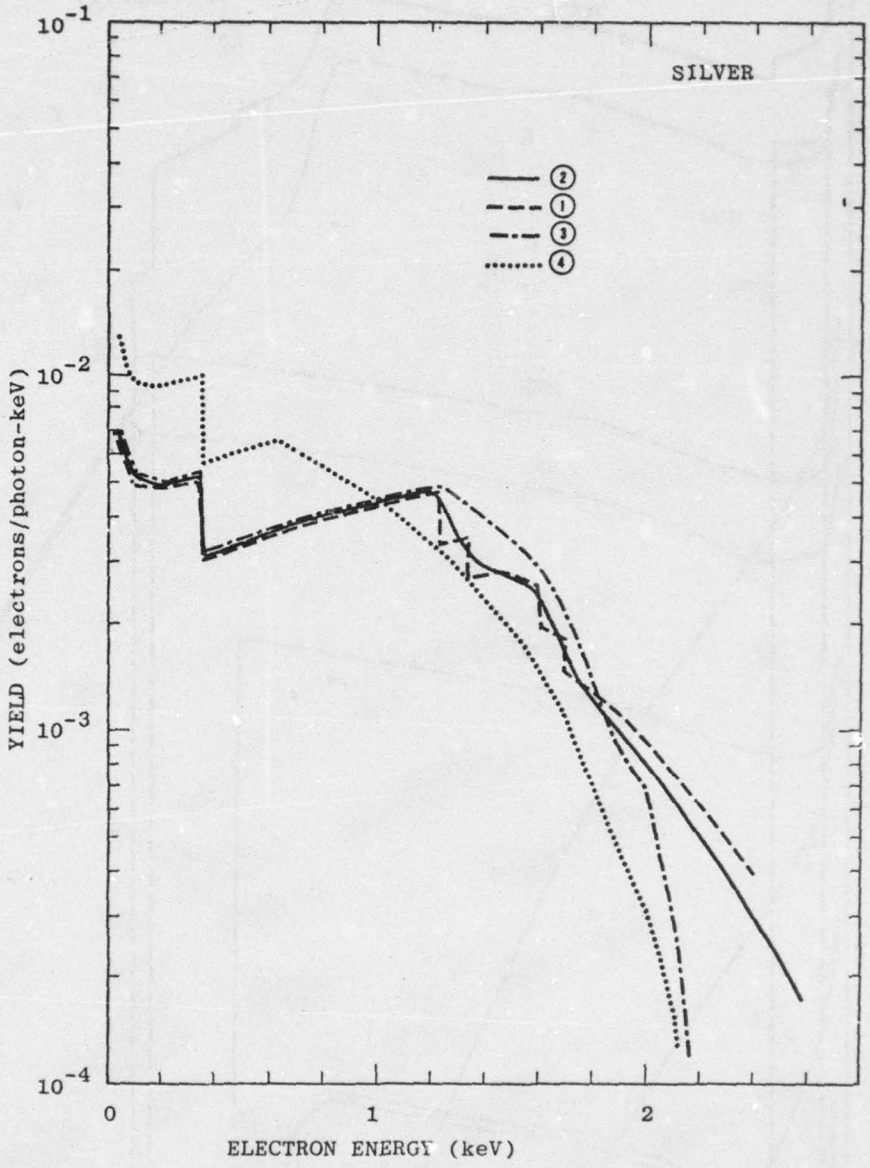

FIGURE 25. Empirical Differential Yields for Ag for Photon Spectra (1)-(4)

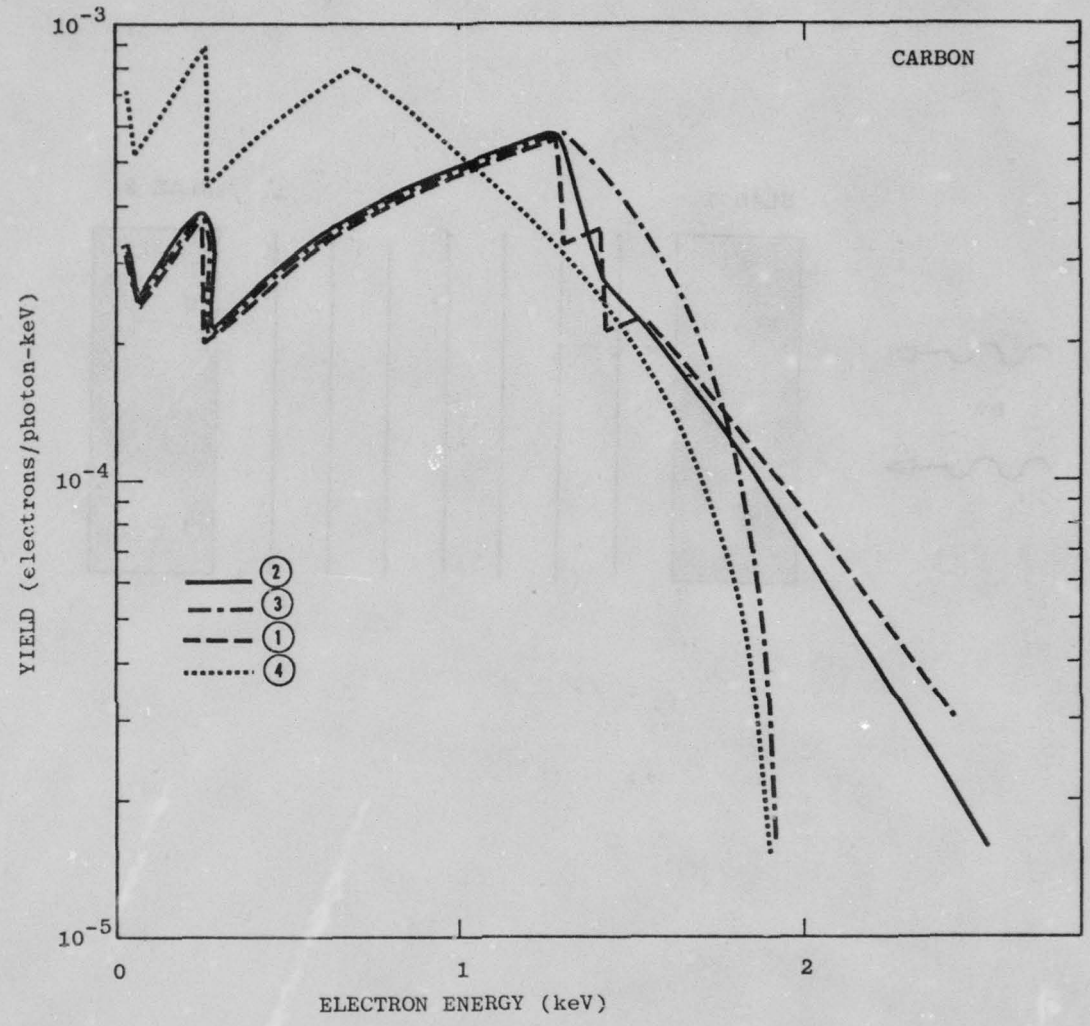

FIGURE 26. Empirical Differential Yields for C for Photon Spectra (1)-(4)

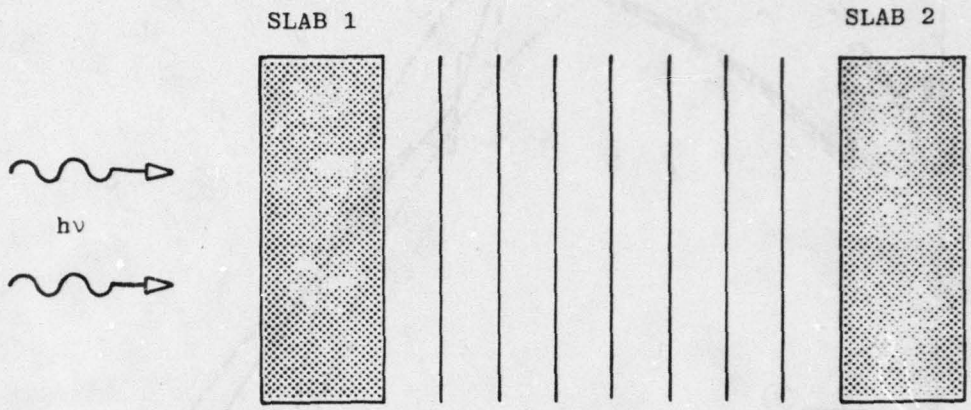

FIGURE 27. Transport Geometry

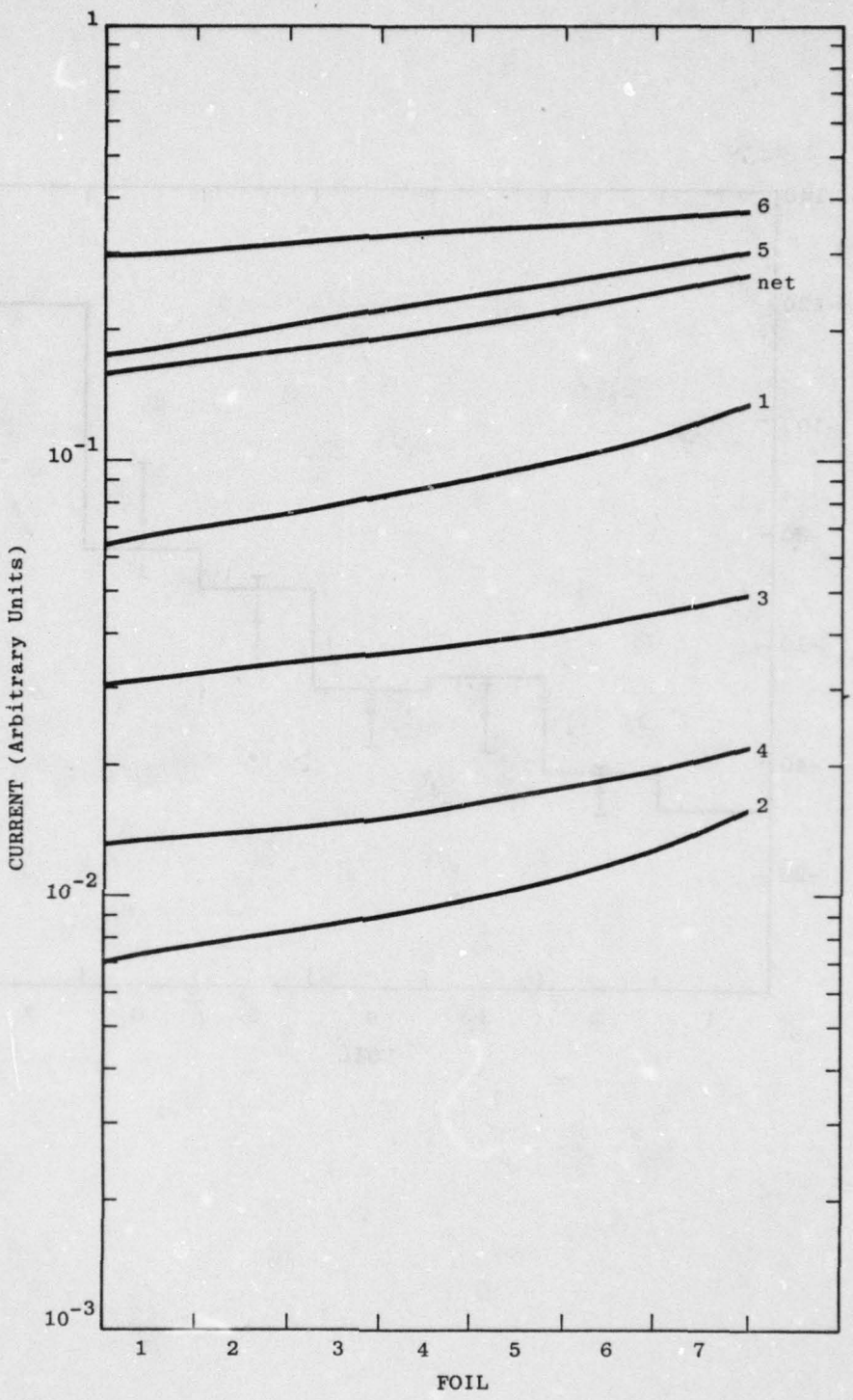

FIGURE 23. Currents in the foils for the Aℓ/Aℓ/Pb Configuration

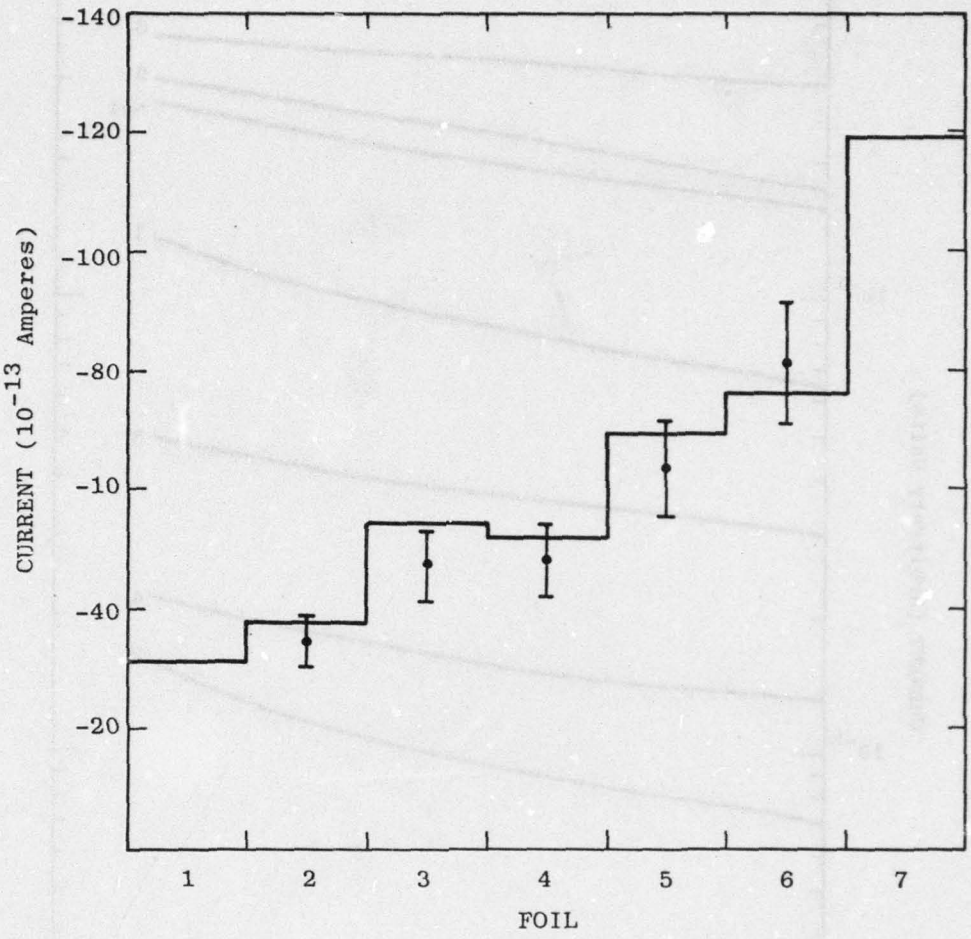

FIGURE 29. Charge Deposition Profile for the Aℓ/Aℓ/Pb Configuration

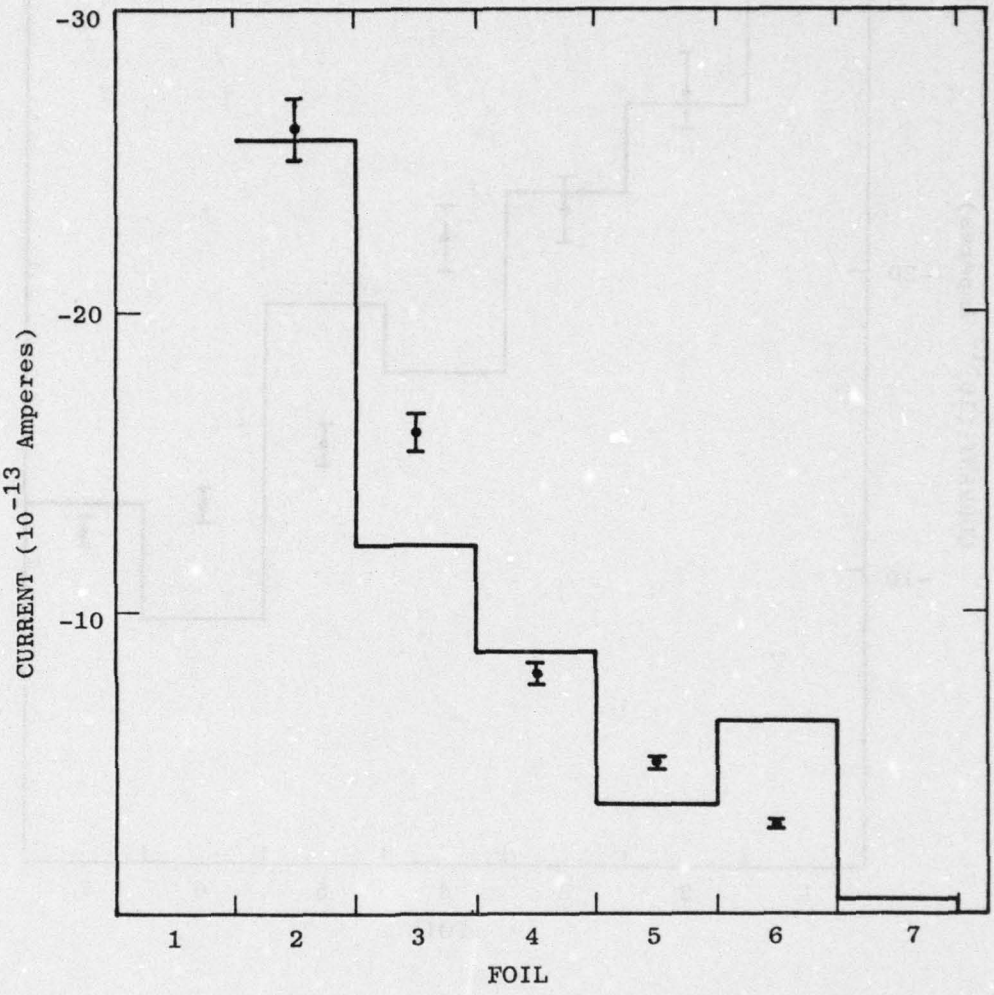

FIGURE 30. Charge Deposition Profile for the Pb/Aℓ/Aℓ Configuration

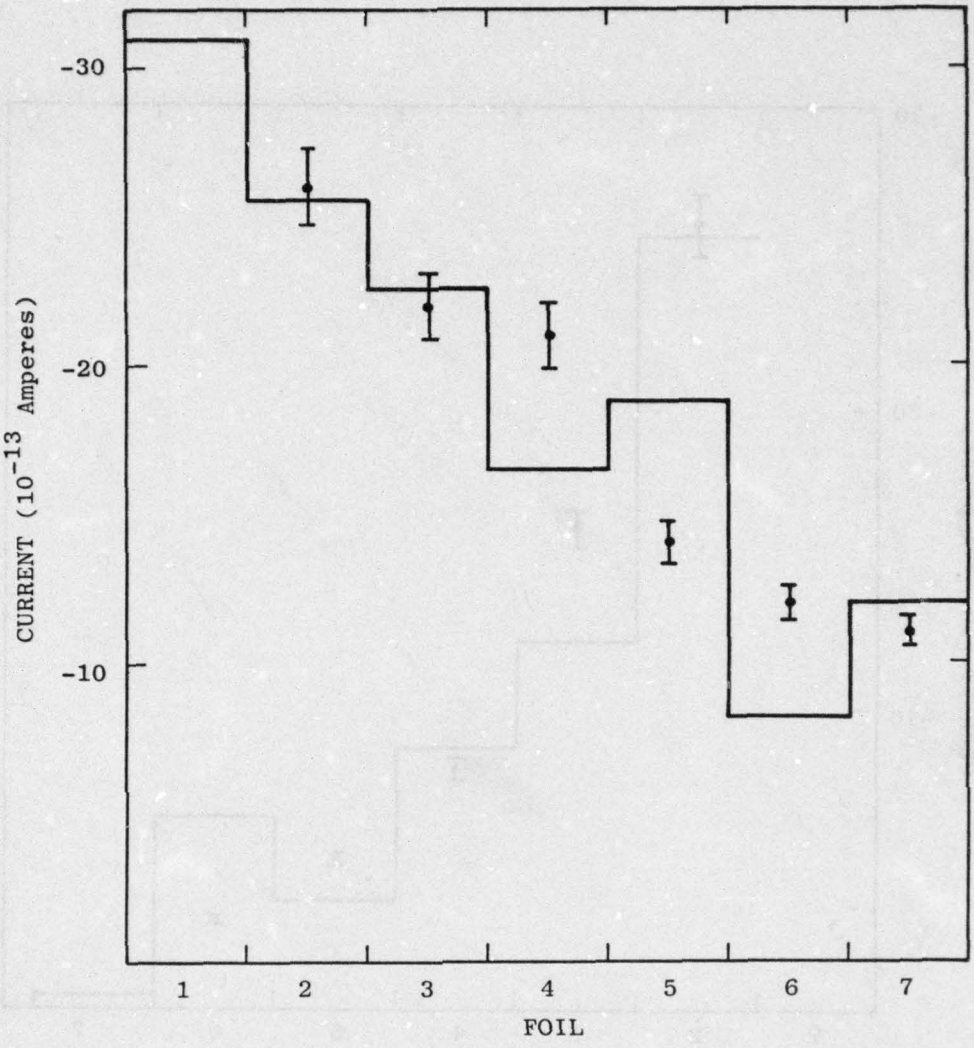

FIGURE 31. Charge Deposition Profile for the Pb/Sn/Sn Configuration

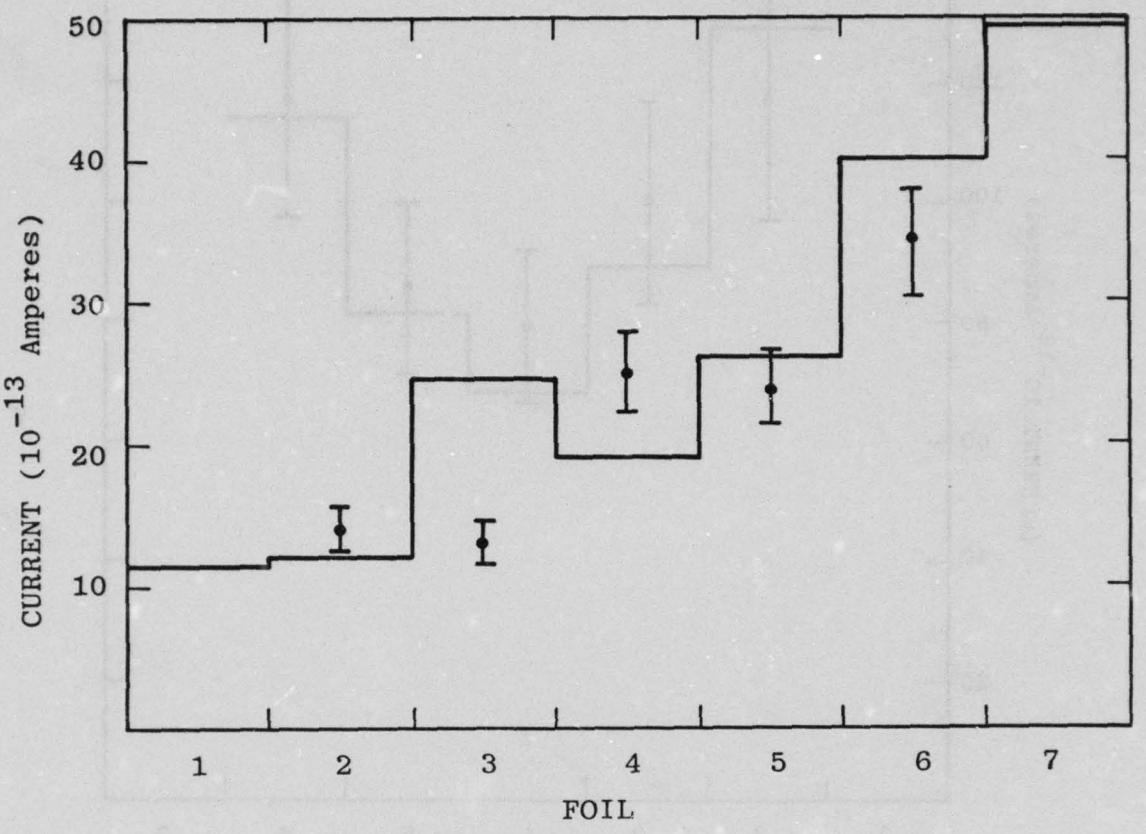

FIGURE 32. Charge Deposition Profile for the Sn/Sn/Aℓ Configuration

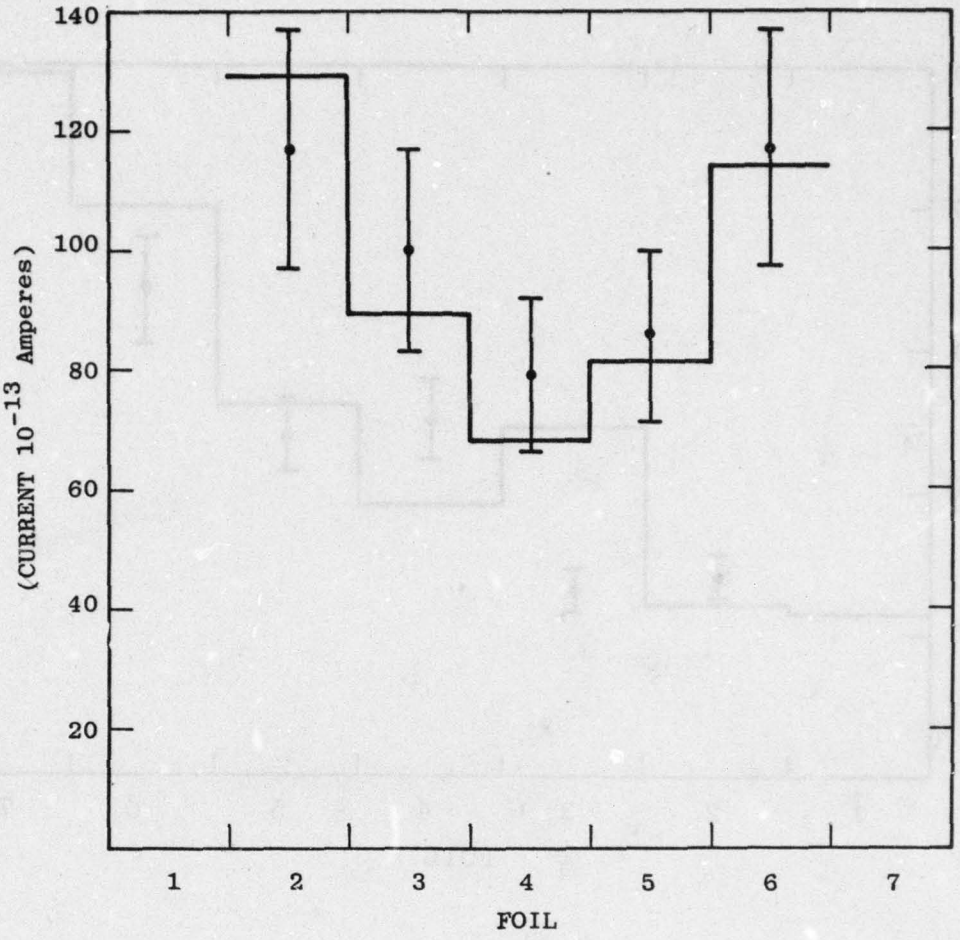

FIGURE 33. Charge Deposition Profile for the Aℓ/Ta/Aℓ Configuration

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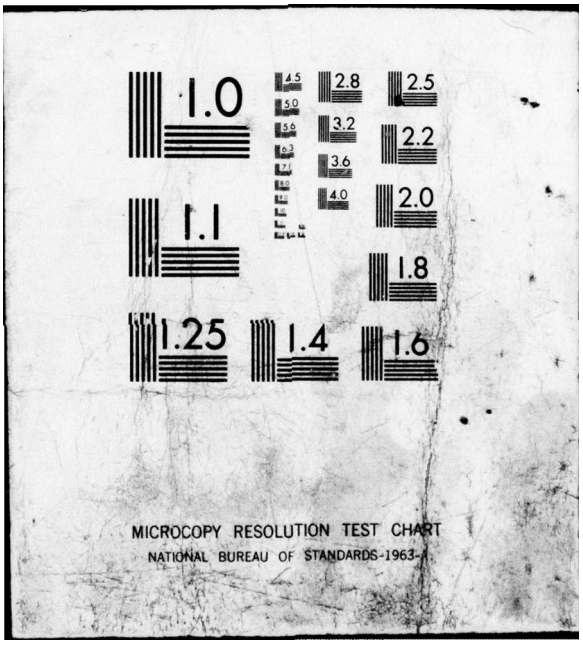
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